

Testing the hypothesis of the Earth's magnetosphere behaving like an avalanching system

A. T. Y. Lui

The Johns Hopkins University Applied Physics Laboratory, Laurel, Maryland 20723, USA

Received: 18 October 2004 – Revised: 9 December 2004 – Accepted: 10 December 2004 – Published: 17 December 2004

Part of Special Issue “Nonlinear processes in solar-terrestrial physics and dynamics of Earth-Ocean-System”

Abstract. The global auroral dissipation power as observed by the imager on the Polar spacecraft is used as a proxy for the power dissipation of the Earth's magnetosphere to examine whether or not the magnetosphere is an avalanching system. It is found that the probability density distributions for the area and power of auroral activity sites have a power law component within a finite scale range, suggestive of a scale-free nature in this finite-size system. This property is robust, prevailing with variations in the threshold used to define auroral activity sites and in the strength of the external driver, namely, the solar wind. The statistical characteristics on the temporal evolution of auroral sites are then examined, which leads to a criterion that can be used to predict about 42 min in advance the total energy dissipation during the lifetime of an auroral activity site. The scale-free characteristics of auroral activity appears to be an intrinsic feature of the magnetosphere based on a comparison of the probability density distribution in the total auroral brightness power with that of the solar wind power input parameters in the same period as the auroral observations. These results are consistent with the hypothesis of the magnetosphere behaving like an avalanching system.

1 Introduction

The Earth's magnetosphere is a voluminous, many-body space system with fascinating natural phenomena invigorating our curiosity. It is often considered as a prototype system consisting of physical processes akin to those responsible for many astrophysical phenomena out of reach by in situ measurements.

A visible manifestation on the disturbance state of the Earth's magnetosphere is the auroral displays seen in the polar regions. Explosive intensifications of auroral displays signify the occurrence of magnetospheric substorms (Aka-

sofu, 1964), which is a natural phenomenon involving disturbances throughout the Earth's immediate space environment (Akasofu, 1968). Research on this subject has now revealed that a proper understanding of magnetospheric activity, such as substorms, requires probing the tight coupling among many plasma regions in the magnetosphere and ionosphere. Furthermore, a substantial number of space disturbances span over a wide range in temporal and spatial scales. These characteristics draw striking similarities with many complex systems in nature (Bak et al., 1987; Lui, 1991; Chang, 1992, 1999; Consolini et al., 1996; Consolini, 1997a,b; Chapman et al., 1998; Angelopolous et al., 1999; Klimas et al., 2000; Uritsky et al., 2002).

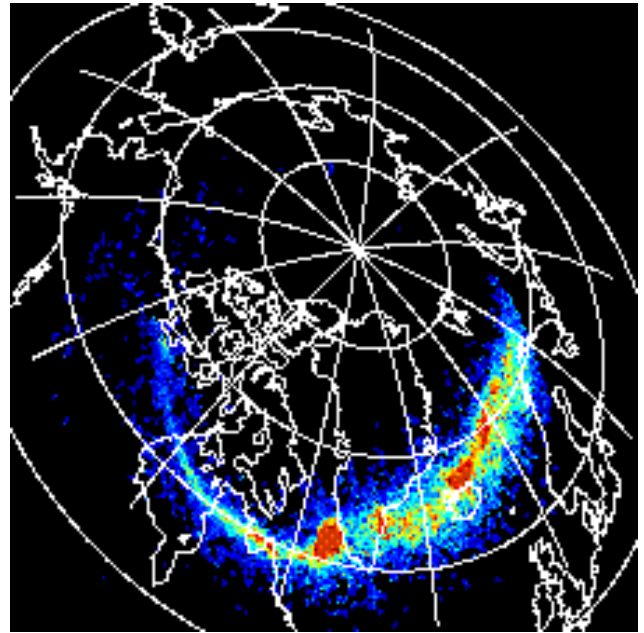
In this paper, we review briefly some findings from studies testing the hypothesis of the magnetosphere being an avalanching system. These results were obtained through investigations on auroral activity. An implicit assumption of these studies is that the energy release associated with magnetospheric disturbances is represented well by the energy dissipated in auroral activity. A criterion formulated to provide forecasting of auroral activity based on specific statistical properties of auroral activity sites is also discussed. After these brief reviews of previously published results that are based on the author's own involvement (Sects. 2 and 3), some new results are presented in Sect. 4 to shed light on the possibility of the scale-free characteristics in auroral activity being merely a result of the scale-free characteristics in the solar wind impinging on the magnetosphere. This question is addressed by examining both the solar wind and the auroral disturbances in the same time period with extreme statistics.

2 Dissipation power and auroral area

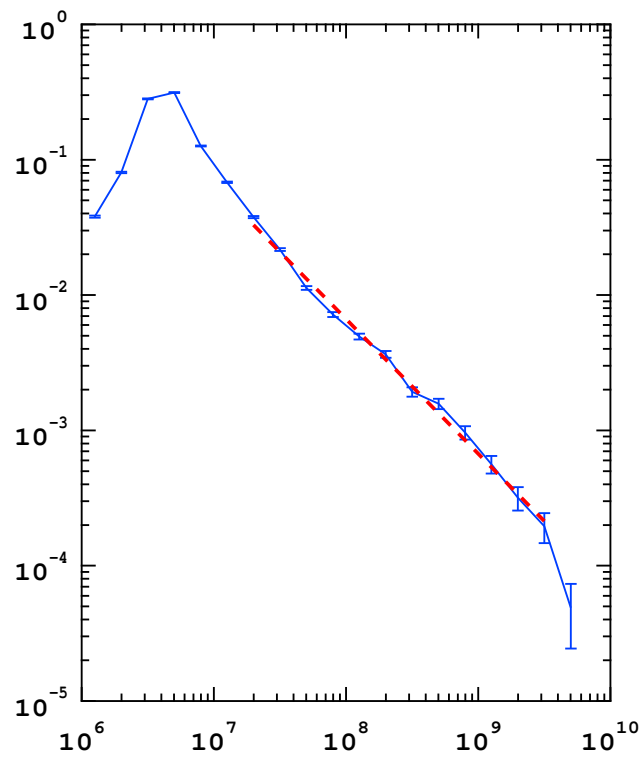
One approach to investigate the analogy between the magnetosphere and an avalanche system is through the distribution of auroral activity area and dissipation power (Lui et al., 2000). The essence and sample result of this approach are summarized in Fig. 1. The schematic in Fig. 1a illustrates



(a)



(b)



(c)

Fig. 1. Auroral power exhibits a scale-free probability density distribution similar to an avalanching system: (a) a schematic of a snow avalanche and magnetotail activity sites responsible for observed auroral dissipation, (b) an example of a global auroral observation from Polar UVI, and (c) a representative distribution of auroral area showing a power law dependence suggestive of its scale-free nature.

this analogy. Localized magnetotail disturbances, which are perceived as the source for auroral activities, may trigger large scale magnetospheric disturbances such as substorms, much like snow avalanches may be triggered by small collapses of snow on a mountain. Figure 1b shows an example of a global auroral image from the Polar UVI instrument. Geographic latitudes and longitudes as well as outline of landmasses are overlaid on the image. Lui et al. (2000) analyzed a month of these global auroral images. The analysis used an auroral intensity threshold significantly above the noise level ($3.5 \text{ erg/cm}^2/\text{s}$) to isolate individual auroral activity sites. For non-substorm intervals, the probability density distribution (PDF) or normalized histogram of the auroral dissipation power (and area also) of these sites was found to have a power law component within a sizeable range of scales. This result is illustrated in Fig. 1c. The normalized occurrence frequency, equivalent to PDF, is plotted as a function of the auroral dissipation power in a log-log graph. A straight line for its slope in this plot implies the occurrence frequency having a power law dependence, i.e. scale-free, on the dissipation power of the activity site. This statistical property on the dissipation power and area of auroral activity sites is consistent with the hypothesis that the magnetosphere behaves like an avalanching system. For substorm intervals, there is an additional component in these distributions, exhibiting a peak at a characteristic scale. These are robust features since they are not affected by variation in the intensity threshold used for the identification of auroral activity sites or the strength of the solar wind energy input (external driver) (Lui, 2002). The observed power and area distributions seem to resemble the PDF of energy dissipation in sandpile avalanches (Chapman et al., 1998). The conclusion reached based on these two previous studies was later supported by a more extensive analysis of auroral images by Uritsky et al. (2002). They combined the spatial and temporal variations of auroral activity sites to enlarge the range of scales under investigation and found the PDF to have a power law shape, revealing the scale-free nature of the phenomenon.

3 Forecast of auroral dissipation

In extending the previous investigations, Lui et al. (2003) tracked the time development of each activity site through the temporal sequence of auroral images. The following procedure was used to identify the continuation of an auroral site in consecutive global auroral images. It is based on overlapping of the site location. If two or more activity sites merge, the site with the largest area is designated as the continuation of the combined site while the others are considered as being terminated. If a single site splits into two or more sites, the site with the largest area is designated as the continuation of the previous site and the others are assigned as new activity sites.

Figures 2a–d shows the statistical properties of four attributes of the auroral activity sites obtained by Lui et al. (2003). These attributes are the total area, instantaneous

dissipation power, the rate of change in area, and the rate of change in dissipation power. The data were sorted into two distributions based on the total energy dissipation using a level of 1 TJ as the boundary dividing the two regimes of low- and high-dissipation. For the six-month period carried out in the analysis, there are 7659 and 1397 activity sites for high- and low-dissipation, respectively. An examination of these results indicates that there are significant differences between these two populations. More specifically, in terms of area (Fig. 2a), low dissipation sites are found below 10^6 km^2 , peaking in their PDF at $10^{4.5} \text{ km}^2$, while the high dissipation sites have the peak for their distribution at $10^{5.5} \text{ km}^2$. A similar remarkable difference is seen in dissipation power (Fig. 2b). Low dissipation sites have dissipation powers below $10^{9.8} \text{ Watt}$, peaking in their distribution at $10^{8.1} \text{ Watt}$, while high dissipation sites have a peak in their distribution at $10^{9.1} \text{ Watt}$. In terms of area change (Fig. 2c), low dissipation sites have mainly values within $\pm 10^3 \text{ km}^2/\text{s}$ while high dissipation sites have a larger spread, peaking at $1500 \text{ km}^2/\text{s}$. The power changes for high- and low-dissipation sites show similar differences (Fig. 2d). These significant differences between high- and low-dissipation sites are not unique to the chosen dissipation threshold since they are also found with lower or higher total energy dissipation thresholds. The remarkable differences between the two populations for these four attributes provided the starting point for Lui et al. (2003) to develop a criterion to separate them.

A systematic search of levels imposed on these four attributes of an activity site led to a criterion suitable to differentiate the high- and low-dissipation sites at a given level of total energy dissipation. For example, the criterion for the prediction of a total energy dissipation of 1 TJ is that the auroral activity site satisfies at least one of the following conditions: area $> 10^6 \text{ km}^2$, power $> 10^{9.2} \text{ Watt}$, area change $> 5000 \text{ km}^2/\text{s}$, and power change $> 10^7 \text{ Watt/s}$. These conditions are to be checked continuously as the auroral activity site evolves in time. The prediction result is given in Fig. 2e, sorting the distributions of low- and high-dissipation predictions in terms of the actual dissipation observed. The high- and low-dissipation domains are shaded with light pink and light blue backgrounds, respectively. The distributions for low and high dissipation sites are plotted with blue and red lines, respectively. There are 9056 activity sites in this study. The number of correct prediction for high dissipation is 918, correct prediction for low dissipation is 7574, incorrect prediction for high dissipation is 404, and incorrect prediction for low dissipation is 160. These give the probabilities for correct high prediction and for correct low prediction to be 0.85 and 0.95, respectively. The false alarm ratio is 0.31. The true skill score for the above criterion is 0.80. Similar favorable results are obtained for prediction of 10 and 0.1 TJ (Lui et al., 2003). From the statistics on the lifetime of each auroral activity site and the time of prediction based on the criterion, it is found that the prediction generally is made on the average of about 42 min in advance before the total energy dissipation of an auroral site during its lifetime can be determined from observation.

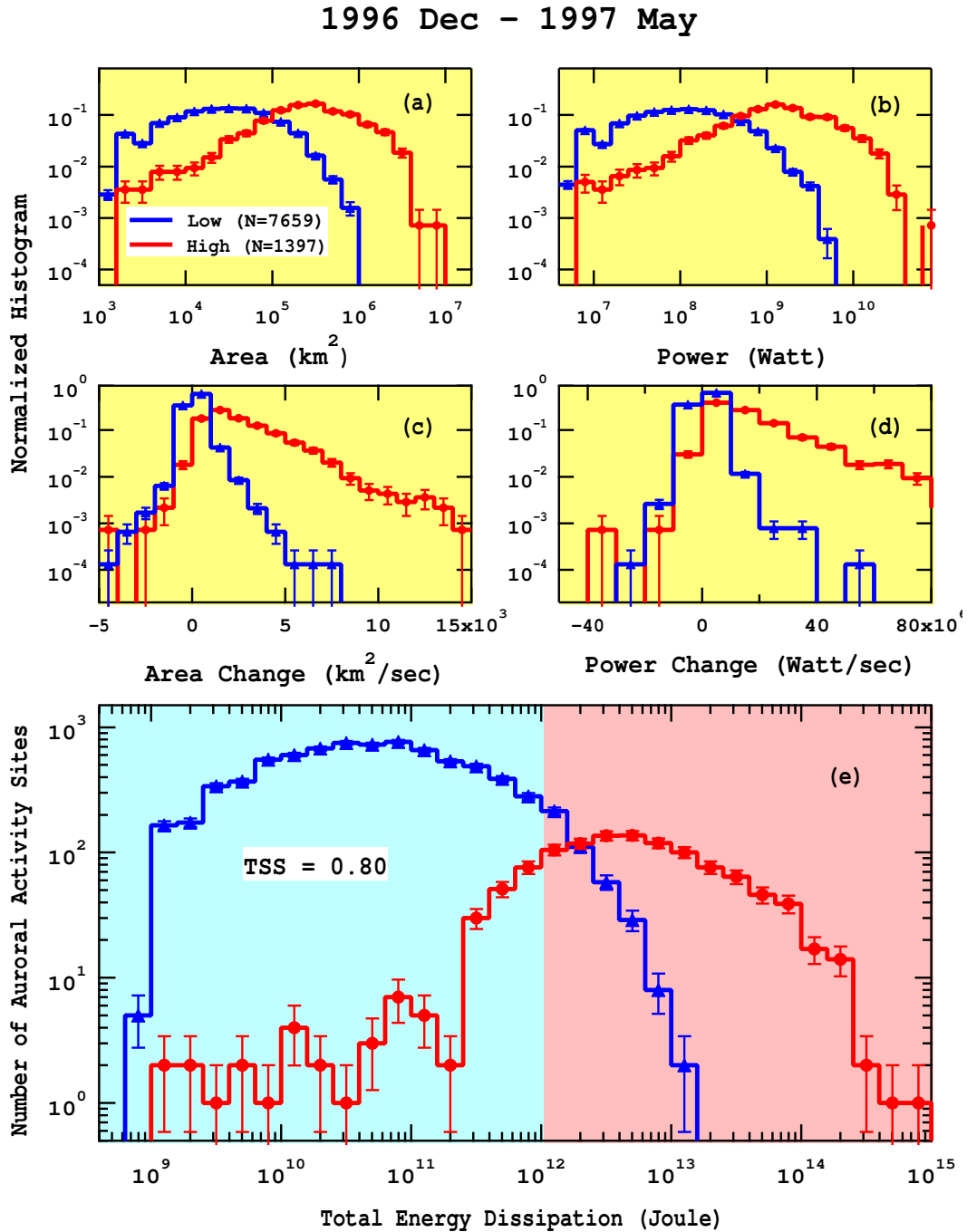


Fig. 2. Statistical results from analysis on the temporal evolution of auroral sites separated by the total energy dissipation threshold of 1 TJ: distribution in (a) area, (b) power, (c) area change, (d) power change, and (e) the prediction of total energy dissipation during the lifetime of an auroral activity site versus the actual observed total energy dissipation.

4 Cause of scale-free property

The hypothesis of magnetosphere behaving like an avalanching system is not universally accepted. For example, Watanabe et al. (2000) examined the AL index and found that the number of occurrence versus the disturbance magnitude has a power law form only when the disturbance is below a certain threshold in magnitude. This is taken as supporting evidence

for the existence of a characteristic time scale for magnetospheric disturbances. Another result casting doubt to the hypothesis is reported by Freeman et al. (2000). They noted that the power law PDFs of burst size and lifetime occur not only in self-organized criticality (SOC) but also in MHD turbulence. The power law feature in burst lifetime distributions actually demonstrates self-similarity or fractality, rather than distinguishes between the physical mechanisms that produce

it. They compared the burst lifetime PDFs of the solar wind power input parameters for the time period of 1995–1998 with those of the AU and AL indices for another time period of 1978–1988. A threshold at the median value of each data set was used in the analysis. More specifically, the solar wind power input parameters used in their study are vB_s , where v is the solar wind speed and B_s is the southward component of the interplanetary magnetic field (IMF), and the episilon parameter (Perreault and Akasofu, 1978), which is given by $v(B_y^2+B_z^2)\sin^4(\theta/2)$, where B_y , B_z are the y - and z -components of the IMF, and θ is the polar angle of the IMF projected on the yz -plane. Their result shows that all these PDFs have power law components with nearly the same value of the power law index. They tentatively interpreted the result to indicate the solar wind turbulence driving directly the magnetospheric disturbance and the magnetosphere may not be an avalanching system.

A sensitive way to determine the causal relationship between the scaling in the solar wind and that in the magnetosphere is to examine the extreme values these parameters exhibit. If the scale-free characteristic of the magnetosphere is inherited from that of the solar wind, it is anticipated that the extreme values in the deviation from the mean value would have identical shapes in their PDFs. This procedure needs to examine both plasma media for the same period of time to be meaningful. In extreme statistics, two types of PDFs are often discussed. One is the Gumbel distribution, which describes the distribution for a global quantity x by

$$G(x)=k(e^{y-e^y})^a, \quad \text{with } y=b(x-s),$$

where k , b , and s are determined by the zeroth, first, and second moments of $G(x)$ ($M0$, $M1$, and $M2$) such that $M0=1$, $M1=0$, and $M2=1$, and a ranges from 1 to 2 (Chapman et al., 2002). The other one is the Frechet distribution with the form similar to Gumbel, i.e.

$$F(x)=k(e^{y-e^y})^a, \quad \text{with } y=s+b \ln(1+x/c(q)),$$

where $c(q)$ is a characteristic value of x and the normalized source PDF $N(x)$ has the form

$$N(x)=N_0/(1+x^2)^q,$$

where N_0 is a constant. Several investigations have indicated that these distributions may represent a universal non-Gaussian fluctuation PDF in highly correlated systems (Bramwell et al., 1998, 2000; Pinton et al., 1999; Zheng and Trimper, 2001).

This new study involves a statistical compilation of the total auroral brightness power over the entire auroral region based on Polar UVI images. No intensity threshold is needed to isolate individual auroral activity sites and no tracking of temporal development of the auroral brightness power over the entire auroral region is needed. An implicit assumption here is that the total auroral brightness power at any moment in time is dominated by the auroral power of the largest activity site. Global auroral images in the six-month period of December 1996–May 1997 were analyzed in this manner.

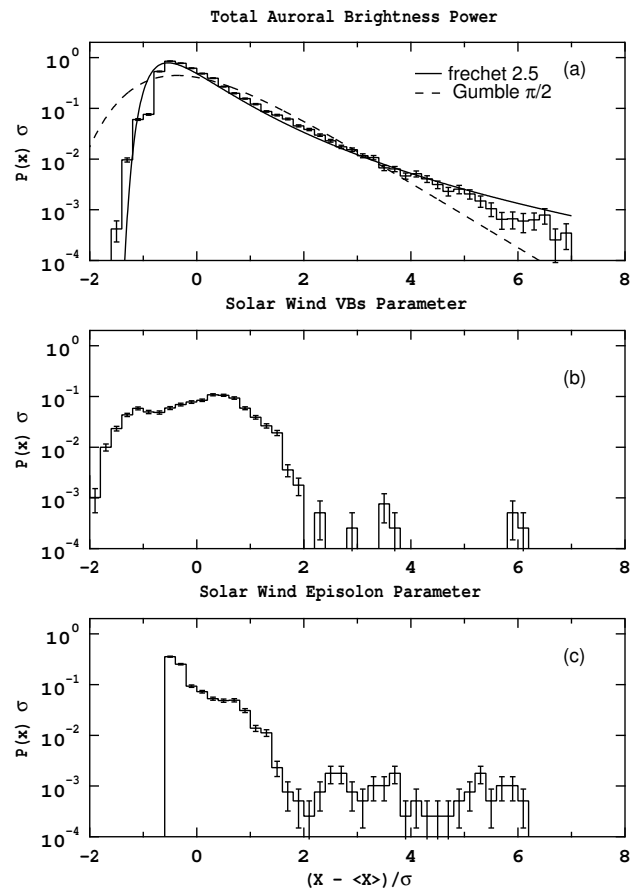


Fig. 3. The normalized distribution weighted by the variance of its dependent variable for (a) the total global auroral brightness at an instant in time, (b) the solar wind power input parameter vB_s , and (c) the solar wind power input parameter episilon.

The solar wind parameters for the corresponding time period were obtained by IMP-8. Two parameters representing the input power from the solar wind to the magnetosphere are vB_s and episilon, as discussed earlier. These parameters were constructed from IMP-8 solar wind data at a time resolution of 1 min.

Figure 3a shows the PDF, weighted by the variance of its dependent variable, of the total auroral brightness power for the six-month period in this study. This distribution seems to be fitted well by the Frechet distribution with $q=2.5$, but not by the Gumbel distribution, which is shown in the figure by the dashed curve for $a=\pi/2$ as a reference. On the other hand, the PDFs of both the solar wind vB_s and episilon parameters given in Figs. 3b and 3c show very distinct shapes. Crudely speaking, the PDF of the solar wind vB_s parameter resembles closer to the Gumbel distribution than to the Frechet distribution, while that of the solar wind episilon parameter resembles closer to the Frechet distribution than to the Gumbel distribution. However, there are large discrepancies of the observed PDFs from these theoretical distributions in the mid-range values of $(X - \langle X \rangle)/\sigma=2-5$ that the observed PDFs are more accurately judged to be quite differ-

ent from either the Gumbel distribution or the Frechet distribution. This result suggests that the scale-free nature in the magnetosphere is not a direct result of the scale-free characteristics of the solar wind.

5 Summary and conclusion

We have presented a brief review of attempts to test the hypothesis of the magnetosphere behaving like an avalanching system by using global auroral dissipation as a proxy for magnetospheric dissipation. The distributions in area and power of auroral activity sites are found to exhibit a power law component suggestive of a scale-free nature. This property appears to be quite robust as it is preserved under different levels of threshold used in the determination of auroral activity sites above the background luminosity as well as under different strength of the external driver. The second study reviewed here involves using the statistical characteristics of these auroral activity sites to develop a criterion to forecast the total auroral energy dissipation during the lifetime of an auroral activity site. The derived criterion appears to be quite accurate with a true skill score of 0.80. The third study presented here tackles the question whether or not the scale-free nature in magnetospheric disturbances is a direct result of the scale-free characteristics of the solar wind power input. This involves extreme statistics in comparing the probability density distribution of the total auroral brightness power over the entire polar region with that of the solar wind power input parameters, namely, the vB_s and the ϵ parameters. The dissimilarity of these distributions favors the suggestion that the scale-free characteristics of the magnetospheric disturbances are intrinsic and not being imposed or transmitted by the external driver.

The above results are consistent with the idea that the magnetosphere is a forced and/or self organized critical system as theorized by a number of researchers (Chang, 1992, 1999; Consolini et al., 1996; Consolini, 1997a, b; Chapman et al., 1998; Angelopoulos et al., 1999; Klimas et al., 2000; Uritsky et al., 2002). However, magnetospheric disturbances can be equally considered as phase transitions for an out-of-equilibrium system, as discussed by several researchers (Consolini and Lui, 2000; Sitnov et al., 2000, 2001).

Acknowledgement. This work is supported by the Atmospheric Sciences Division of the NSF Grant ATM-0135667 and by the NASA Grant NNG04G128G to The Johns Hopkins University Applied Physics Laboratory.

Edited by: A. Chian

Reviewed by: two referees

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