

Charged black holes in Vaidya backgrounds: Hawking's Radiation

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Abstract

In this paper we propose a class of embedded solutions of Einstein's field equations describing non-rotating Reissner-Nordstrom-Vaidya and rotating Kerr-Newman-Vaidya black holes. The Reissner-Nordstrom-Vaidya is obtained by embedding Reissner-Nordstrom solution into non-rotating Vaidya. Similarly, we also find the Kerr-Newman-Vaidya black hole, when Kerr-Newman embeds into the rotating Vaidya solution. The Reissner-Nordstrom-Vaidya solution is type D whereas Kerr-Newman-Vaidya metric is algebraically special type II of Petrov classification of space-time. These embedded solutions can be expressed in Kerr-Schild ansatz on different backgrounds. The energy momentum tensors for both non-rotating as well as rotating embedded solutions satisfy the energy conservation equations which show that they are solutions of Einstein's field equations. The surface gravity, area, temperature and entropy are also presented for each embedded black hole. It is observed that the area of the embedded black holes is greater than the sum of the areas of the individual ones. By considering the charge to be a function of radial coordinate it is shown that there is a change in the masses of the variable charged black holes. If such radiation continues, the mass of the black hole will evaporate completely thereby forming 'instantaneous' charged black holes and creating embedded 'negative mass naked singularities' describing the possible life style of radiating embedded black holes during their continuous radiation processes.

Keywords: Reissner-Nordstrom-Vaidya; Kerr-Newman-Vaidya black hole; rotating Vaidya; surface gravity; Hawking's radiation.

1 Introduction

In the theory of black hole expressed in spherical polar coordinates, there is a singularity at the origin $r = 0$; whereas at $r = \infty$, the metric approaches the Minkowski flat space. The event horizons of a black hole are expressed by this coordinate. Thus, the nature of black holes depend on the radial coordinate r . In an earlier paper [1] it is shown that Hawking's radiation [2] can be expressed in classical space-time metrics, by considering the charge e to be function of the radial coordinate r of non-rotating Reissner-Nordstrom as well as rotating Kerr-Newman black holes. The variable-charge $e(r)$ with respect to the coordinate r is followed from Boulware's suggestion [3] that *the stress-energy tensor may be used*

to calculate the change in the mass due to the radiation. Accordingly, the energy-momentum tensor of a particular space can be used to calculate the change in the mass in order to incorporate the Hawking's radiation effects in classical space-time metrics. This idea suggests to consider the stress-energy tensor of electromagnetic field of different forms or functions from those of Reissner-Nordstrom, as well as Kerr-Newman, black holes as these two black holes do not seem to have any direct Hawking's radiation effects on the change of masses in space-time metrics. Thus, a variable charge in the field equations will have the different function of the energy-momentum tensor of a charged black hole. *Such a variable charge e with respect to the coordinate r in Einstein's equations is referred to as an electrical radiation (or Hawking's electrical radiation) of a black hole* [1]. For every electrical radiation the charge e is considered to be a function of r in analyzing the Einstein-Maxwell field equations and it has shown *mathematically* how the electrical radiation induces to produce the changes of the masses of *variable*-charged black holes. The idea of losing (or changing) mass can be incorporated in the space-time metrics at the rate as the electrical energy is radiated from the charged black hole. In fact, the change in the mass of a charged black hole takes place due to the vanishing of Ricci scalar of the electromagnetic field. It was found that every electrical radiation $e(r)$ of the black hole leads to a reduction in its mass by some quantity. If we consider such electrical radiation taking place continuously for a long time, then a continuous reduction of the mass will take place in the black hole body, and the original mass of the black hole will evaporate completely. At that stage the complete evaporation of the mass will lead the gravity of the object depending only on the electromagnetic field, and not on the mass of black hole. Such an object with zero mass is being referred as an 'instantaneous' naked singularity *i.e.* a naked singularity that exists for an instant and then continues its electrical radiation to create negative mass [1]. So this naked singularity is different from the one mentioned in Steinmular *et al* [4], Tipler *et al* [5] in the sense that an 'instantaneous' naked singularity, discussed in [4, 5] exists only for an instant and then disappears.

It is also noted that the time taken between two consecutive radiations is supposed to be so short that one may not physically realize how quickly radiations take place. Thus, it seems natural to expect the existence of an 'instantaneous' naked singularity with zero mass only for an instant before continuing its next radiation to create a negative mass naked singularity. This suggests that it may also be possible in the common theory of black holes that, as a black hole is invisible in nature, one may not know whether, in the universe, a particular black hole has mass or not, but electrical radiation may be detected on the black hole surface. Immediately after the complete evaporation of the mass, if one continues to radiate the remaining remnant, there may be a formation of a new mass. If one repeats the electrical radiation further, the new mass might increase gradually and then the space-time geometry will represent the negative mass naked singularity. The classical space-time metrics, for both stationary rotating and non-rotating, which represent the

negative mass naked singularities have given in [1].

The black hole embedded into other spaces play an important role in relativistic framework. For example, the cosmological constant is found present in the inflationary scenario of the early universe in a stage where the universe is geometrically similar to the original de Sitter space [6]. Also black holes embedded into de Sitter cosmological universe can avoid the direct formation of negative mass naked singularities during Hawking's electrical radiation process of black holes [7]. It is also known that the rotating Vaidya black hole with variable mass is a non-stationary solution of Einstein's field equations and is non-vacuum *algebraically special* of Petrov classification of space-time [8, 9, 10]. In fact the non-stationary Vaidya black holes is commonly known with two different characters – one is non-rotating [11] and other rotating [8, 9, 10]. This is due the fact that there are two stationary black holes vacuum solutions – one is non-rotating Schwarzschild and other rotating Kerr. Hence it is highly necessary to understand the Vaidya black holes in two classes – one is *non-rotating* and other *rotating*. This is because if we wish Vaidya black hole to embed into non-rotating Reissner-Nordstrom black hole, we need the non-rotating Vaidya black hole for similar configuration; whereas for rotating Kerr-Newman we should have rotating Vaidya black hole for the same rotating system.

The aim of this paper is to study Hawking's radiation of Reissner-Nordstrom and Kerr-Newman black hole embedded into the Vaidya backgrounds. The solutions derived here describe non-rotating Reissner-Nordstrom black hole embedded into the non-rotating Vaidya to obtain a Reissner-Nordstrom-Vaidya and rotating embedded black hole *i.e.*, the Kerr-Newman black hole is embedded into the rotating Vaidya null radiating space generating the Kerr-Newman-Vaidya. The definitions of these embedded black holes are in agreement with the one defined by Cai *et.al.* [12], that *when the Schwarzschild black hole is embedded into the de Sitter space, one has the Schwarzschild-de Sitter black hole*. Thus, these solutions describe *non-rotating*, Reissner-Nordstrom-Vaidya and rotating, Kerr-Newman-Vaidya black holes. These embedded solutions can be expressed in Kerr-Schild forms to regard them as the extension of Glass-Krisch superposition [13], which is further the extension of that of Xanthopoulos [14]. These Kerr-Schild ansatze show that these embedded black holes are solutions of Einstein's field equations. It is also noted that such generation of embedded solutions in *non-rotating* cases may be seen in [15].

In an introductory survey Hawking and Israel [16] have discussed the black hole radiation in three possible ways with suggestive remarks –‘So far there is no good theoretical frame work with which to treat the final stages of a black hole but there seem to be three possibilities: (i) The black hole might disappear completely, leaving just the thermal radiation that it emitted during its evaporation. (ii) It might leave behind a non-radiating black hole of about the Planck mass. (iii) The emission of energy might continue indefinitely creating a negative mass naked singularity’. In order to examine Hawking's radiation in classical space-time metrics, the suggestion of Boulware [3] leads to consider the stress-energy tensors of electromagnetic

field of different forms or functions from those of Reissner-Nordstrom, as well as Kerr-Newman, black holes as these two black holes do not seem to have any direct Hawking's radiation effects. Considering the charge $e(r)$ to be a function of radial coordinate r , *i.e.* having different structure of the source of energy-momentum tensors, Hawking electrical radiation of Reissner-Nordstrom and Kerr-Newman black holes has been developed in the forms of space-time metrics in [1]. As a result, it appears that (i) the changes in the mass of black holes, (ii) the formation of 'instantaneous' naked singularities with zero mass and (iii) the creation of 'negative mass naked singularities' in Reissner-Nordstrom as well as Kerr-Newman black holes may presumably be the correct formulations in classical space-time metrics of the three possibilities of black hole evaporation suggested by Hawking and Israel [16]. It has been shown that, during the complete evaporation of the masses of both Reissner-Nordstrom and Kerr-Newman black holes, the electrical radiation continues indefinitely creating negative mass naked singularities. The same work of electrical radiation process has been extended in the presence of de Sitter cosmological background as Reissner-Nordstrom-de Sitter and Kerr-Newman-de Sitter in [7]. Due to the presence of the de Sitter background, the direct formation of 'negative mass naked singularities' could be avoided by creating "instantaneous" charged de Sitter cosmological universe during radiation process. Here in this paper we are interested to develop a similar work of Hawking's electrical radiation on a different background having a variable mass function, rather than a constant on de Sitter background; *i.e.* Hawking's radiation of black holes on Vaidya background.

The paper is organized as follows: In Sections 2 and 3, we develop embedded Reissner-Northrom-Vaidya and Kerr-Newman-Vaidya black holes by using Wang-Wu mass functions in Newman-Penrose spin coefficient formalism [17]. We discuss the physical properties of the solutions observing the nature of their energy momentum tensors and Weyl scalars. We also present the surface gravity, entropy and angular velocity for each of these embedded black holes as they are important parameters of a black hole. Sections 4 and 5 deal with the Hawking's radiation on the variable-charged Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black holes. We present various classical space-time metrics affected by the change in the masses describing the possible life style of radiating embedded black holes in different stages during radiation process. In section 6 we conclude with a discussion of the results. The relativistic aspect of Hawking's radiation has been studied in these embedded black holes. In Appendix A, for a better explanation of the paper, we present an energy-momentum tensor for a metric having mass function $\hat{M}(u, r)$ of two variables. We also present the general equations for the energy conservation laws associated with $\hat{M}(u, r)$ in Newman-Penrose formalism. Appendix B is devoted for the NP quantities of the embedded Kerr-Newman-Vaidya solution and has shown that the energy-momentum tensors for the embedded black holes obey the energy conservation laws. The content of the paper are summarized in the form of theorems as:

Theorem 1. *Embedded Reissner-Nordstrom-Vaidya is a Petrov type D in the classification of space-times, whose one of the repeated principal directions ℓ_a is geodesic, expanding, shear free and zero twist.*

Theorem 2. *Embedded Kerr-Newman-Vaidya is an algebraically special type II in the Petrov classification of space-times, whose one of the repeated principal directions ℓ_a is geodesic, expanding, shear free and non-zero twist.*

Theorem 3. *The horizon areas of the embedded Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya are greater than the sum of the horizon areas of individual black holes, or the entropies of the embedded black holes are greater than the sum of the entropies of the individual ones.*

Theorem 4. *Embedded Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black holes can be expressed in Kerr-Schild ansatz in different backgrounds.*

Theorem 5. *Every electrical radiation of Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black holes will lead to a change in their masses without affecting the Maxwell scalar as well as Vaidya mass.*

Theorem 6. *The non-rotating and rotating charged Vaidya metrics describe instantaneous naked singularities during the Hawking's evaporation process of electrical radiation of Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black holes.*

Theorem 7. *During the radiation process, after the complete evaporation of masses of both variable-charged Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black holes, the electrical radiation will continue indefinitely creating negative mass naked singularities on Vaidya background.*

Theorem 8. *If an electrically radiating black hole, rotating or non-rotating, is embedded into Vaidya spaces, it will continue to embed into the same space forever.*

We find that Theorem 3 is consistent with the suggestion of Bekenstein [18] that “when two black holes merge, the area of the resulting black hole (provided, of course, that one forms) cannot be smaller than the sum of the initial areas”. The idea of initial black holes comes from the embedded ones: in the case of embedded Reissner-Nordstrom-Vaidya, the Reissner-Nordstrom and the non-rotating Vaidya are regarded as initial (or individual) ones, and similarly Kerr-Newman and rotating Vaidya are the initial (or individual) ones in Kerr-Newman-Vaidya. Theorem 4 shows that embedded Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black holes are solutions of Einstein's field equations. It is found that the Theorems 5, 6 and 7 are in favour of the three possibilities of black holes made by Hawking and Israel [16], and indicate the various stages of the life of embedded radiating black holes. Theorem 7 provides a violation of Penrose's cosmic-censorship hypothesis that “no naked singularity can ever be created” [19]. Theorem 8 shows the nature of

embedded radiating black holes that continue to embed into the same space. Here it is convenient to use the phrase ‘change in the mass’ rather than ‘loss of mass’ as there may be a possibility of creation of mass after the exhaustion of the original mass if one continues the same process of electrical radiation. This will be seen later in the paper. The presentation of this paper is essentially based on the differential form technique developed by McIntosh and Hickman [20] in (-2) signature.

2 Reissner-Nordstrom-Vaidya Solution

In this section for deriving an embedded Reissner-Nordstrom-Vaidya we consider a general canonical metric in Eddington-Finkelstein coordinate system given in (A.2) of Appendix A, with the quantities μ^* , ρ^* and p (when the rotational parameter a sets to zero) in (A.5). However, there is no straightforward way to derive an embedded Reissner-Nordstrom-Vaidya solution by solving the non-linear Einstein’s field equations associated with the line element (A.2) having a variable mass function $M(u, r)$. Therefore, in order to embed the Reissner-Nordstrom into Vaidya space, we choose, without loss of generality, the $q_n(u)$ in the Wang-Wu mass function $M(u, r)$ [15] as follows

$$q_n(u) = \begin{cases} M + f(u), & \text{when } n = 0 \\ -e^2/2, & \text{when } n = -1 \\ 0, & \text{when } n \neq 0, -1, \end{cases} \quad (2.1)$$

where M and e are constants. Then, the mass function takes the form

$$\hat{M}(u, r) = \sum_{n=-\infty}^{+\infty} q_n(u) r^n = M + f(u) - \frac{e^2}{2r}. \quad (2.2)$$

Using this mass function in general canonical metric (A.2) in Eddington-Finkelstein coordinate system, we find a line element

$$ds^2 = [1 - r^{-2}\{2r(M + f(u)) - e^2\}] du^2 + 2du dr - r^2 d\theta^2 - r^2 \sin^2\theta d\phi^2. \quad (2.3)$$

The constants M and e are related with the mass and charge of Reissner-Nordstrom solution respectively. $f(u)$ is the variable mass of Vaidya solution, u is the retarded time which is used for calculating the outgoing or radiating energy-momentum tensor. When $f(u)$ sets to zero, the above line element will reduce to the Reissner-Nordstrom solution. The choice of mass function in (2.2) is acceptable in the framework of general relativity as the linear superposition of particular known solutions in $M(u, r)$ is also a solution of Einstein’s field equations.

Now by using (2.) in (A.3-A.5) when $a = 0$, we find the energy-momentum tensor for the Reissner-Nordstrom-Vaidya solution as

$$T_{ab} = \mu^* \ell_a \ell_b + 2\rho^* \ell_{(a} n_{b)} + 2p m_{(a} \bar{m}_{b)} \quad (2.4)$$

which can be expressed in two parts as

$$T_{ab}^{(n)} = \mu^* \ell_a \ell_b, \quad (2.5)$$

$$T_{ab}^{(E)} = 2\rho^* \ell_{(a} n_{b)} + 2p m_{(a} \bar{m}_{b)}, \quad (2.6)$$

where the null density μ^* , the energy density ρ^* and pressure p are found as follows

$$\mu^* = -\frac{2}{K r^2} f(u)_{,u}, \quad \rho^* = p = \frac{e^2}{K r^4}. \quad (2.7)$$

Here ℓ_a and n_a are real, and m_a is complex null vectors with the normalization conditions $\ell_a n^a = 1 = -m_a \bar{m}^a$; and they are given by

$$\begin{aligned} \ell_a &= \delta_a^1, \quad n_a = \left[1 - r^{-2} \left\{ 2r(M + f(u)) - e^2 \right\} \right] \delta_a^1 + \delta_a^2, \\ m_a &= -\frac{r}{\sqrt{2}} \left\{ \delta_a^3 + i \sin \theta \delta_a^4 \right\}. \end{aligned} \quad (2.8)$$

We use the notation $2\ell_{(a} n_{b)} = \ell_a n_b + n_a \ell_b$. $T_{ab}^{(n)}$ is the energy-momentum tensor for the Vaidya null radiating fluid and $T_{ab}^{(E)}$ for non-null electromagnetic field existing in non-rotating Reissner-Nordstrom-Vaidya metric g_{ab} . The energy-momentum tensor (2.4) can be regarded as the electromagnetic field $T_{ab}^{(E)}$ is interacting with the null radiating fluid $T_{ab}^{(n)}$, and is traceless ($T = 0$). T^{ab} in (2.4) obeys the energy conservation equations

$$T^{ab}{}_{;b} = 0. \quad (2.9)$$

showing that the non-rotating Reissner-Nordstrom-Vaidya metric is an exact solution of Einstein's field equation.

We also find the Weyl scalar for the solution (2.3) as follows

$$\psi_2 = -\frac{1}{r^4} \{ r(M + f(u)) - e^2 \}. \quad (2.10)$$

The non-vanishing of Weyl scalar ψ_2 means that the Reissner-Nordstrom-Vaidya solution is Petrov type D whose one of the repeated principal directions ℓ_a is geodesic ($\kappa = \epsilon = 0$), shear free ($\sigma = 0$) and expanding ($\hat{\theta} \neq 0$) with zero twist ($\hat{\omega} = 0$). (Here κ, ϵ, σ , etc. are Newman-Penrose spin coefficients for the line element (2.3) and can be obtained from (B.2) by setting rotational parameter $a = 0$.) This completes the proof of the Theorem 1.

The invariants of Ricci and Riemann Curvature tensors for the solution are given as follows

$$R_{ab} R^{ab} = \frac{4}{r^8} e^4 \quad (2.11a)$$

$$R_{abcd} R^{abcd} = \frac{1}{r^8} \left[48 \left(r \{ M + f(u) \} - e^2 \right)^2 + 20e^4 \right]. \quad (2.11b)$$

These invariants become infinite when $r \rightarrow 0$. $K \rightarrow \infty$, which suggests that the Reissner-Nordstrom-Vaidya is a model of a primordial black hole [21].

The solution of Reissner-Nordstrom-Vaidya will describe a black hole with external horizon r_+ and internal Cauchy horizon r_- given by

$$r_{\pm} = \{M + f(u)\} \pm \sqrt{\{M + f(u)\}^2 - e^2}, \quad (2.12)$$

when $\Delta \equiv r^2 - 2r\{M + f(u)\} - e^2 = 0$. The areas at these horizons are obtained as

$$\begin{aligned} A_{\pm}^{\text{RNV}} &= \int_0^{\pi} \int_0^{2\pi} \sqrt{g_{\theta\theta}g_{\phi\phi}} d\theta d\phi \Big|_{r=r_{\pm}} \\ &= 8\pi \left\{ (M + f(u))^2 - \frac{e^2}{2} \pm (M + f(u))\sqrt{(M + f(u))^2 - e^2} \right\}. \end{aligned} \quad (2.13)$$

It is observed that when $f(u)$ is set to zero, the area A_{\pm}^{RNV} will reduce to that of Reissner-Nordstrom black hole, whereas if we set M and e to be zero, the area A_{\pm}^{RNV} (2.13) will be that of Vaidya black hole as

$$A_{\pm}^{\text{RN}} = 8\pi \left\{ M^2 - \frac{e^2}{2} \pm M\sqrt{M^2 - e^2} \right\}, \quad (2.14a)$$

$$A_+^{\text{V}} = 16\pi f^2(u) \quad \text{and} \quad A_-^{\text{V}} = 0. \quad (2.14b)$$

From (2.13) and (2.14) we find that the area A_{\pm}^{RNV} of the embedded Reissner-Nordstrom-Vaidya black hole is greater than the sum of the areas of the individual ones *i.e*

$$A_{\pm}^{\text{RNV}} > A_{\pm}^{\text{RN}} + A_{\pm}^{\text{V}}. \quad (2.15)$$

which proves Theorem 3 for non-rotating case stated above in the introduction. This greater area theorem of the embedded black hole is in agreement with Bekenstein's suggestion [18]. Now from the entropy-area relation $S = A/4$ [22], we obtain the entropies of the horizons

$$S_{\pm}^{\text{RNV}} = 2\pi \left[\{M + f(u)\}r_{\pm} - \frac{e^2}{2} \right], \quad (2.16)$$

which satisfy the equality relations $S_{\pm}^{\text{RNV}} > S_{\pm}^{\text{RN}} + S_{\pm}^{\text{V}}$. Here S_{\pm}^{RN} and S_{\pm}^{V} denote the entropies of Reissner-Nordstrom and Vaidya black holes respectively.

The surface gravity κ of a horizon is defined by the relation $n^b \nabla_b n^a = \kappa n^a$, where the null vector n^a in (2.8) above is parameterized by the coordinate u , such that $d/du = n^b \nabla_b$ [23, 24]. The surface gravities of the horizons at $r = r_+$ and $r = r_-$ are found as

$$\mathcal{K}_+ = -\frac{1}{2r_+^2}(r_+ - r_-), \quad \mathcal{K}_- = \frac{1}{2r_-^2}(r_+ - r_-) \quad (2.17)$$

and their ratio is

$$\mathcal{K} = -\frac{r_+^2}{r_-^2} \quad (2.18)$$

Hawking's temperature at horizons r_+ and r_- are given by

$$T_+ = \frac{\mathcal{K}_+}{2\pi} = -\frac{1}{4\pi r_+^2}(r_+ - r_-), \quad T_- = \frac{\mathcal{K}_-}{2\pi} = \frac{1}{4\pi r_-^2}(r_+ - r_-). \quad (2.19)$$

Since there are two horizons, there must be two different temperatures given by T_+ and T_- from the event and Cauchy horizon respectively. It is also observed that the ratio of the temperatures at the horizons is same as the ratio of the surface gravities.

The non-rotating Reissner-Nordstrom-Vaidya solution can be expressed in Kerr-Schild ansatz on the Reissner-Nordstrom background g_{ab}^{RN} as

$$g_{ab}^{RNV} = g_{ab}^{RN} - \{2f(u)/r\}\ell_a\ell_b. \quad (2.20)$$

Also the metric can be expressed on Vaidya background g_{ab}^V as

$$g_{ab}^{RNV} = g_{ab}^V - \{2M/r - e^2/r^2\}\ell_a\ell_b. \quad (2.21)$$

The expressibility of these ansatz on different background spaces, means that it is always true to talk about either Reissner-Nordstrom black hole embedded into the Vaidya space as Reissner-Nordstrom-Vaidya or the Vaidya black hole embedded into the Reissner-Nordstrom space to form Vaidya-Reissner-Nordstrom – both possess the same geometrical meaning. This indicates that in the nature of the embedded black hole, one cannot physically predict which black hole starts first to embed into another. It is also found that the two metric tensors of Reissner-Nordstrom and Vaidya solutions cannot be added to emerge into an embedded Reissner-Nordstrom-Vaidya solution as

$$g_{ab}^{RNV} \neq \frac{1}{2}(g_{ab}^{RN} + g_{ab}^V);$$

however the energy-momentum tensors of the respective solutions can be added to obtain that of the embedded one. It is to mention that the Kerr-Schild ansatz (2.20) and (2.21) are the generalizations of Xanthopoulos [14] as well as Glass and Krisch [13] in the charged Reissner-Nordstrom solution. These ansatz shows the proof of the non-rotating part of Theorem 2 above.

3 Kerr-Newman-Vaidya Solution

In this section, we shall develop an embedded Kerr-Newman-Vaidya black hole which will be the generalization of the non-rotating Reissner-Nordstrom-Vaidya solution (2.3), by using the rotating metric (A.1) of Appendix A with a mass function $M(u, r)$ and the quantities μ^* , ρ^* , p and ω^* in (A.5). However, there is no straightforward way of solving the Einstein's field equations associated with the line element (A.1) for obtaining a Kerr-Newman-Vaidya solution. Hence, in order to obtain the

rotating solution, we utilize the mass function (2.2) in the rotating frame (A.1) given in Appendix A and find a line element describing Kerr-Newman-Vaidya solution as

$$\begin{aligned}
ds^2 = & [1 - R^{-2}\{2r(M + f(u)) - e^2\}] du^2 + 2du dr + 2aR^{-2}\{2r(M + f(u)) \\
& - e^2\} \sin^2\theta du d\phi - 2a \sin^2\theta dr d\phi - R^2 d\theta^2 \\
& - \{(r^2 + a^2)^2 - \Delta a^2 \sin^2\theta\} R^{-2} \sin^2\theta d\phi^2,
\end{aligned} \tag{3.1}$$

where $\Delta = r^2 - 2r\{M + f(u)\} + a^2 + e^2$. Here M and e are the mass and the charge of Kerr-Newman solution, a is the rotational parameter and $f(u)$ represents the mass function of rotating Vaidya null radiating fluid. When we set $f(u) = 0$, the metric (3.1) recovers the usual Kerr-Newman black hole, and if $M = 0$, then it is the ‘rotating’ charged Vaidya null radiating black hole [8, 9, 10]. The line element (3.1) describes the embedded Kerr-Newman-Vaidya black hole having singularity when $\Delta = 0$.

In this rotating solution, the Vaidya null fluid is interacting with the non-null electromagnetic field F_{ab} , whose Maxwell scalar is defined by $\phi_1 = \frac{1}{2}F_{ab}(\ell^a n^b + \bar{m}^a m^b)$. Then, by using (2.2) in (A.3) we find an energy-momentum tensor for the rotating line element (3.1) as follows:

$$T_{ab} = \mu^* \ell_a \ell_b + 2\omega \ell_{(a} \bar{m}_{b)} + 2\bar{\omega} \ell_{(a} m_{b)} + 2\rho^* \ell_{(a} n_{b)} + 2p m_{(a} \bar{m}_{b)}. \tag{3.2}$$

Here ℓ_a and n_a are real and m_a is complex null vectors with the normalization conditions $\ell_a n^a = 1 = -m_a \bar{m}^a$, and are given in (B.1). This T_{ab} may be expressed in terms of those of the electromagnetic field as well as the rotating Vaidya null fluid respectively as

$$T_{ab}^{(E)} = 2\rho^* \ell_{(a} n_{b)} + 2p m_{(a} \bar{m}_{b)} \tag{3.3}$$

$$T_{ab}^{(n)} = \mu^* \ell_a \ell_b + 2\omega \ell_{(a} \bar{m}_{b)} + 2\bar{\omega} \ell_{(a} m_{b)}, \tag{3.4}$$

where the quantities are

$$\rho^* = p = \frac{e^2}{K R^2 R^2}, \tag{3.5a}$$

$$\mu^* = -\frac{r}{K R^2 R^2} \left\{ 2r f(u)_{,u} + a^2 \sin^2\theta f(u)_{,uu} \right\}, \tag{3.5b}$$

$$\omega = -\frac{ia \sin\theta}{\sqrt{2} K \bar{R} R^2} f(u)_{,u}. \tag{3.5c}$$

Here ρ^* and p are denoted respectively the energy density and pressure of the electromagnetic field of Kerr-Newman black hole, and μ^* being the null density contributed from the Vaidya mass $f(u)$. ω is the rotational density arisen from the parameter a , that the null fluid $T_{ab}^{(n)}$ is rotating as the expression of ω involves the rotating parameter a coupling with $\partial f(u)/\partial u$; both are non-zero quantities for the existence of the rotating Vaidya null radiating solution. It is to mention that the energy-momentum tensor (3.2) obeys the energy conservation equations

$$T^{ab}{}_{;b} = 0. \tag{3.6}$$

as shown in (B.7) of Appendix B. This indicates that the metric tensor of Kerr-Newman-Vaidya associated with the line element (3.1) is a solution of Einstein's field equations.

When $\Delta = 0$, the solution (3.1) will describe an embedded black hole with external event horizon r_+ and internal Cauchy horizon r_- which are found as

$$r_{\pm} = \{M + f(u)\} \pm \sqrt{[M + f(u)]^2 - a^2 - e^2}, \quad (3.7)$$

and the non-stationary limit surface $g_{uu} = 0$ is given by

$$r \equiv r_e(u, \theta) = \{M + f(u)\} + \sqrt{[M + f(u)]^2 - a^2 \cos^2 \theta - e^2}. \quad (3.8)$$

which coincides with the event horizon r_+ only at the points $\theta = 0$ and $\theta = \pi$. This feature distinguishes the Kerr-Newman-Vaidya space-time from the Reissner-Nordstrom-Vaidya one. This situation cannot be observed in non-rotating solution. The area of the horizon is found as

$$\begin{aligned} A_{\pm}^{\text{KNV}} &= \int_0^{\pi} \int_0^{2\pi} \sqrt{g_{\theta\theta}g_{\phi\phi}} d\theta d\phi \Big|_{r=r_{\pm}} \\ &= 4\pi \left\{ (M + f(u))^2 - e^2 \pm 2(M + f(u)) \sqrt{(M + f(u))^2 - a^2 - e^2} \right\}. \end{aligned} \quad (3.9)$$

When $f(u)$ is set to zero, the area A_{\pm}^{KNV} will reduce to that of Kerr-Newman black hole A_{\pm}^{KN} ; whereas if we set M and e to be zero the area will be that of the rotating Vaidya black holes A_{\pm}^{V} as

$$A_{\pm}^{\text{KN}} = 4\pi \left\{ 2M^2 - e^2 \pm 2M \sqrt{M^2 - a^2 - e^2} \right\} \quad (3.10a)$$

$$A_{\pm}^{\text{V}} = 4\pi \left\{ 2f^2(u) \pm 2f(u) \sqrt{f^2(u) - a^2} \right\}. \quad (3.10b)$$

We observe the difference between the areas of the rotating (3.9) and non-rotating (2.14b) Vaidya solutions, that the non-rotating Vaidya does not have A_-^{V} . From (3.9) and (3.10), we find that the area of embedded black hole is greater than the sum of the areas of the individual black holes as

$$A_{\pm}^{\text{KNV}} > A_{\pm}^{\text{KN}} + A_{\pm}^{\text{V}} \quad (3.11)$$

which establishes the proof of Theorem 3 for the rotating part. The inequality relation (3.11) is consistent with Bekenstein's suggestion [18] that, *if two black holes emerges to form a black hole, the area of the later will be greater than the sum of the individual ones*. Now, from the entropy-area relation of Bekenstein and Hawking [22], we obtain the entropy

$$S_{\pm}^{\text{KNV}} = 2\pi \left[\{M + f(u)\} r_{\pm} - \frac{e^2}{2} \right]. \quad (3.12)$$

This entropy satisfies the inequality relation $S_{\pm}^{\text{KNV}} > S_{\pm}^{\text{KN}} + S_{\pm}^{\text{V}}$, which is true for the rotating and non-rotating embedded black holes. It is also found that the sum of the entropies at the event (r_+) and Cauchy (r_-) horizons does not affect by the rotational parameter a as $S = S_{r_+}^{\text{KNV}} + S_{r_-}^{\text{KNV}} = 4\pi(M + f(u))^2 - 2\pi e^2$.

The surface gravity of the horizons at $r = r_{\pm}$ are

$$\mathcal{K}_+ = \frac{1}{2R_+^2}(r_- - r_+), \quad \mathcal{K}_- = \frac{1}{2R_-^2}(r_+ - r_-) \quad (3.13)$$

where $R_{\pm}^2 = r_{\pm}^2 + a^2 \cos^2 \theta$. The ratio of the two surface gravities, \mathcal{K} is

$$\mathcal{K} = \frac{\mathcal{K}_+}{\mathcal{K}_-} = -\frac{R_-^2}{R_+^2}. \quad (3.14)$$

According to Padmanabhan and Roy Choudhury [25], there is a global temperature if the ratio is rational. Hawking's temperatures associated with \mathcal{K}_+ , and \mathcal{K}_- are obtained as

$$T_+ = \frac{\mathcal{K}_+}{2\pi} = \frac{r_- - r_+}{4\pi R_+^2}, \quad T_- = \frac{\mathcal{K}_-}{2\pi} = -\frac{r_- - r_+}{4\pi R_-^2}. \quad (3.15)$$

This shows that there are two different temperatures – one flowing from the event horizon while other from the Cauchy horizon. It is also observed that the ratio of the temperatures is equal to the ratio of the surface gravities (3.14). The angular velocity of the horizons take the form

$$\Omega_{\pm} = \lim_{r \rightarrow r_{\pm}} \left(-\frac{g_{u\phi}}{g_{\phi\phi}} \right) = \frac{a[2r\{M + f(u)\} - e^2]}{(r^2 + a^2)^2} \Big|_{r=r_{\pm}}. \quad (3.16)$$

This will vanish when the rotational parameter a sets to zero. This shows the fact that the non-rotating solution does not have any angular velocity. For instance, the Reissner-Nordstrom-Vaidya black hole discussed above has no angular velocity.

The rotating Kerr-Newman-Vaidya metric (3.1) can be expressed in Kerr-Schild form on the Kerr-Newman background as

$$g_{ab}^{\text{KNV}} = g_{ab}^{\text{KN}} + 2Q(u, r, \theta)\ell_a\ell_b \quad (3.17)$$

where $Q(u, r, \theta) = -rf(u)R^{-2}$, and the vector ℓ_a is a geodesic, shear free, expanding as well as rotating null vector of both g_{ab}^{KN} as well as g_{ab}^{KNV} . g_{ab}^{KN} is the Kerr-Newman metric with $m = e = \text{constant}$. This null vector ℓ_a is one of the double repeated principal null vectors of the Weyl tensor of g_{ab}^{KN} . It appears that the rotating Kerr-Newman geometry may be regarded as joining smoothly with the rotating Vaidya geometry at its null radiative boundary, as shown by Glass and Krisch [13] in the case of Schwarzschild geometry joining to the non-rotating Vaidya space-time. The Kerr-Schild form (3.17) will recover that of Xanthopoulos [14] $g'_{ab} = g_{ab} + \ell_a\ell_b$, when $Q(u, r, \theta) \rightarrow 1/2$ and that of Glass and Krisch [13] $g'_{ab} = g_{ab}^{\text{Sch}} - \{2f(u)/r\}\ell_a\ell_b$ when

$e = a = 0$ for non-rotating Schwarzschild background space. Thus, one can consider the Kerr-Schild form (3.17) as the extension of those of Xanthopoulos as well as Glass and Krisch. When we set $a = 0$, this rotating Kerr-Newman-Vaidya solution (3.1) will recover to non-rotating Reissner-Nordstrom-Vaidya solution (2.3) with the Kerr-Schild form $g_{ab}^{\text{RNV}} = g_{ab}^{\text{RN}} - \{2f(u)/r\}\ell_a\ell_b$, which is still a generalization of Xanthopoulos as well as Glass and Krisch in the charged Reissner-Nordstrom solution.

To interpret the Kerr-Newman-Vaidya solution as the Kerr-Newman black hole embedded into the rotating Vaidya background, we can write the Kerr-Schild form (3.17) as

$$g_{ab}^{\text{KNV}} = g_{ab}^{\text{V}} + 2Q(r, \theta)\ell_a\ell_b \quad (3.18)$$

where $Q(r, \theta) = -(rM - e^2/2)R^{-2}$. Here, g_{ab}^{V} is the rotating Vaidya null radiating black hole and ℓ_a is the geodesic null vector for both g_{ab}^{KNV} and g_{ab}^{V} . *i.e.* $g_{ab}^{\text{KNV}}\ell^a\ell^b = 0 = g_{ab}^{\text{V}}\ell^a\ell^b$. When we set $f(u) = a = 0$, g_{ab}^{V} will recover the flat metric, then g_{ab}^{KNV} becomes the original Kerr-Schild form written in spherical symmetric flat background. It is worth to mention the fact that the rotating embedded solution (3.1) cannot be considered as a bimetric theory since

$$g_{ab}^{\text{KNV}} \neq \frac{1}{2}(g_{ab}^{\text{KN}} + g_{ab}^{\text{V}}). \quad (3.19)$$

This indicates the fact that one cannot simply add the two solutions g_{ab}^{KN} and g_{ab}^{V} in order to get embedded solution g_{ab}^{KNV} . However their energy momentum tensors $T_{ab}^{(n)}$ and $T_{ab}^{(E)}$ given in (3.3) and (3.4) respectively can be added to obtain T_{ab} as in (3.2) above.

These two Kerr-Schild forms (3.17) and (3.18) certainly confirm that the metric g_{ab}^{KNV} is a solution of Einstein's field equations since the background rotating metrics g_{ab}^{KN} and g_{ab}^{V} are solutions of Einstein's equations. They both possess different stress-energy tensors $T_{ab}^{(n)}$ and $T_{ab}^{(E)}$ given in (3.3) and (3.4) respectively. Looking at the Kerr-Schild form (3.18), the Kerr-Newman-Vaidya black hole can be treated as a generalization of Kerr-Newman black hole by incorporating Visser's suggestion [26] that *Kerr-Newman black hole embedded in an axisymmetric cloud of matter would be of interest*. The expressibility of this embedded black hole in different ansatz (3.17), (3.18) means that, it is always true to talk about either Kerr-Newman black hole embedded into the Vaidya space as Kerr-Newman-Vaidya or the Vaidya black hole embedded into the Kerr-Newman space to form Vaidya-Kerr-Newman – both nomenclature possess the same geometrical meaning. That is, one cannot physically predict which black hole started first to embed into another black hole. These Kerr-Schild ansatz (3.17) and (3.18) establish the proof of the rotating part of the Theorem 2.

The line element of Kerr-Newman-Vaidya black hole can also be expressed in

the coordinate system (t, x, y, z) as

$$ds^2 = dt^2 - dx^2 - dy^2 - dz^2 - \frac{2r^3}{r^4 + a^2z^2} \left\{ M + f(u) - \frac{e^2}{2r} \right\} \times \left[dt - \frac{1}{r^2 + a^2} \left\{ r(xdx + ydy) + a(xdy - ydx) \right\} - \frac{z}{r} dz \right]^2 \quad (3.20)$$

where r is defined in terms of x, y and z as

$$r^4 - (x^2 + y^2 + z^2 - a^2)r^2 - a^2z^2 = 0$$

with the following transformations

$$\begin{aligned} x &= (r \cos \phi + a \sin \phi) \sin \theta, \\ y &= (r \sin \phi - a \cos \phi) \sin \theta, \\ z &= r \cos \theta, \quad t = u + r. \end{aligned} \quad (3.21)$$

Then, the above transformed metric (3.20) can be written in the Kerr-Schild form on flat background as

$$g_{ab}^{\text{KNV}} = \eta_{ab} + 2H(t, x, y, z)\ell_a\ell_b, \quad (3.22)$$

where η_{ab} is the flat metric and

$$\begin{aligned} H(t, x, y, z) &= -\frac{2r^3}{r^4 + a^2z^2} \left\{ M + f(t, r) - \frac{e^2}{2r} \right\}, \\ \ell_a dx^a &= dt - \frac{1}{(r^2 + a^2)} \left\{ r(xdx + ydy) + a(xdy - ydx) \right\} - \frac{1}{r} z dz. \end{aligned} \quad (3.23)$$

Here ℓ_a is the null vector with respect to g_{ab}^{KNV} and η_{ab} , and is given in null coordinate system (u, r, θ, ϕ) in Appendix A below. The Kerr-Schild form (3.22) shows that the Kerr-Newman-Vaidya solution expressed in coordinate system (3.21) is a solution of Einstein's field equations. The transformed metric (3.20) is analytic everywhere except at $x^2 + y^2 + z^2 = a^2$ and $z = 0$.

4 Changing mass of variable charged Reissner-Nordstrom-Vaidya black hole

In this section, as a part of discussion of the physical properties of the embedded solutions, we shall discuss a scenario which is capable of avoiding the formation of negative mass naked singularity during Hawking radiation process in Reissner-Nordstrom-Vaidya space-time metrics, describing the life style of embedded radiating black hole. The formation of negative mass naked singularities in classical space-time metrics are being shown in [1] after the complete evaporation of the masses of *non-embedded* rotating Kerr-Newman and non-rotating Reissner-Nordstrom, black holes due to Hawking electrical radiation.

Here we shall clarify two similar nomenclatures having different meaning like Hawking's radiation and Vaidya null radiation. Hawking's radiation is continuous radiation process of energy from the black hole body thereby leading to the change in mass [1-5], whereas the Vaidya null radiation means that the stress-energy momentum tensor describing the gravitation in Vaidya space-time metric is a null radiating fluid [11]. So Vaidya null radiation does not have any direct relation with Hawking's radiation of black holes [18]. By electrical radiation of a charged black hole we mean the variation of the charge e with respect to the coordinate r in the stress-energy momentum tensor of electromagnetic field, thereby able to establish the change of mass of the black hole.

As mentioned earlier in the introduction, the variation of the charge e will certainly lead different forms or functions of the stress-energy tensor from that of Reissner-Nordstrom-Vaidya black hole. To observe the change in the mass of black hole in the space-time metric, one has to consider a different form or function of stress-energy tensor of a particular black hole. That is, in order to incorporate the Hawking's radiation in this black hole (2.3), we must have a different stress-energy tensor as the Reissner-Nordstrom-Vaidya black hole with T_{ab} (2.4) does not show any direct Hawking's radiation effects. The consideration of different forms of stress-energy-momentum tensor in the study of Hawking's radiation effect in classical space-time metrics here is in agreement with Boulware's suggestion [3] mentioned in introduction above.

The line element of the variable-charged Reissner-Nordstrom-Vaidya solution with the assumption that the charged e is a function of coordinate r in (2.3), is given by

$$ds^2 = [1 - r^{-2}\{2r(M + f(u)) - e^2(r)\}] du^2 + 2du dr - r^2 d\theta^2 - r^2 \sin^2\theta d\phi^2. \quad (4.1)$$

If we set $e(r) = \text{constant}$ initially, the metric will recover the charged Reissner-Nordstrom-Vaidya solution (2.3). Here with the variable charge $e(r)$ in (4.1) we find the different energy-momentum tensor from that of the Reissner-Nordstrom-Vaidya solution. Accordingly, from the Cartan's second equations, we calculate the tetrad components of Ricci and Weyl tensors and found as follows:

$$\begin{aligned} \phi_{11} &= \frac{1}{2R^2 R^2} \left\{ e^2(r) - 2r e(r) e(r)_{,r} \right\} + \frac{1}{4R^2} \left\{ e^2(r)_{,r} + e(r) e(r)_{,rr} \right\} \\ \phi_{22} &= -\frac{f(u)_{,u}}{r^2} \end{aligned} \quad (4.2)$$

$$\Lambda = -\frac{1}{12r^2} \left\{ e^2(r)_{,r} + e(r) e(r)_{,rr} \right\}, \quad (4.3)$$

$$\begin{aligned} \psi_2 &= \frac{1}{\bar{R} \bar{R} R^2} \left[e^2(r) - R\{M + f(u)\} + \frac{1}{6} \bar{R} \bar{R} \left\{ e^2(r)_{,r} + e(r) e(r)_{,rr} \right\} \right. \\ &\quad \left. - \bar{R} e(r) e(r)_{,r} \right]. \end{aligned} \quad (4.4)$$

Here we find the non-zero Ricci scalar $\Lambda = \frac{1}{24}g^{ab}R_{ab}$ in (4.3). However, for an electromagnetic field, this Ricci scalar Λ has to vanish leading to the equation

$$e^2(r)_{,r} + e(r)e(r)_{,rr} = 0 \quad (4.5)$$

which can be solved having the solution

$$e^2(r) = 2rm_1 + C \quad (4.6)$$

where m_1 and C are real constants. Substituting this $e^2(r)$ in (4.2) and (4.4), and after identifying the constant $C \equiv e^2$, we obtain

$$\phi_{11} = \frac{1}{2r^4}e^2, \quad (4.7)$$

$$\psi_2 = -\frac{1}{r^4}\left[r\{M - m_1 + f(u)\} - e^2\right]. \quad (4.8)$$

and ϕ_{22} is remained unaffected as in (4.2) above. Accordingly, from (4.7) the Maxwell scalar ϕ_1 with constant e becomes

$$\phi_1 = \frac{e}{\sqrt{2}r^2} \quad (4.9)$$

since $\phi_{11} = \phi_1\bar{\phi}_1$ for electromagnetic field. However, from the Weyl scalar (4.8) we have clearly seen a change in the mass M by some constant quantity m_1 for the first electrical radiation in the embedded black hole. Then the total mass of the Reissner-Nordstrom-Vaidya black hole becomes $(M - m_1) + f(u)$ and the metric takes the form

$$ds^2 = [1 - r^{-2}\{2r(M - m_1 + f(u)) - e^2\}] du^2 + 2du dr - r^2 d\theta^2 - r^2 \sin^2\theta d\phi^2. \quad (4.10)$$

Since, for the first radiation, the Maxwell scalar (4.9) remain the same, radiation, we again consider the charge e to be a function of r for the second in the above metric (4.10) which certainly leads, by the Einstein-Maxwell field equations with the vanishing of Λ , to reduce another quantity m_2 (say) from the total mass, *i.e.* after the second radiation the mass becomes $M - m_1 - m_2 + f(u)$. Again, the Maxwell scalar ϕ_1 remains unchanged and also there is no effect on Vaidya mass function $f(u)$. Thus, if we consider n radiations, considering the charge e to be a function of r , the Einstein's field equations will imply that the mass of the embedded black hole will take the form

$$\mathcal{M} = M - (m_1 + m_2 + m_3 + m_4 + \dots + m_n) + f(u) \quad (4.11)$$

without affecting the Vaidya mass function $f(u)$ and the Maxwell scalar ϕ_1 . At some instant, there will be an thermal equilibrium state between M and amount of radiation. If it still continues to radiate the equilibrium state will be disturbed.

Taking Hawking's radiation of charged black hole embedded in the rotating Vaidya null radiating space, one might expect that the mass M may be radiated away, just leaving $M - (m_1 + m_2 + m_3 + m_4 + \dots + m_n)$ equivalent to the Planck mass of about 10^{-5} g and $f(u)$ untouched; that is, M may not be exactly equal to $(m_1 + m_2 + m_3 + m_4 + \dots + m_n)$, but has a difference of about a Planck-size mass. Otherwise, the mass M may be evaporated completely after continuous radiation, when $M = (m_1 + m_2 + m_3 + m_4 + \dots + m_n)$, just leaving the Vaidya mass function $f(u)$ and the electric charge e only. Thus, we can show this situation of the black hole in the form of a classical space-time metric without the mass M as

$$ds^2 = [1 - r^{-2}\{2rf(u) - e^2\}] du^2 + 2du dr - r^2 d\theta^2 - r^2 \sin^2\theta d\phi^2. \quad (4.12)$$

The Weyl scalar of this metric becomes

$$\psi_2 = \frac{1}{r^4} \left\{ e^2 - r f(u) \right\}. \quad (4.13)$$

and ψ_3, ψ_4 are unaffected. We observe that the remaining metric (4.12) is the non-rotating charged Vaidya null radiating black hole with $f(u) > e^2$ for $u = u_0$. The surface gravity of the horizons at

$$r = r_{\pm} = f(u) \pm \sqrt{\{f(u)^2 - e^2\}} \quad (4.14)$$

is $\mathcal{K}_{\pm} = r_{\pm}^{-2} \{r_{\mp} - r_{\pm}\}$. Consequently, the Hawking's temperature and entropy associated with the horizons are $T_{\pm} = (1/4\pi) r_{\pm}^{-2} \{r_{\mp} - r_{\pm}\}$ and $\mathcal{S} = 2\pi f(u) \{f(u) + \sqrt{f(u)^2 - e^2}\} - \pi e^2$. If it goes on radiating, the mass of Reissner-Nordstrom becomes zero. But the presence of the Vaidya mass function $f(u)$ can avoid the formation of an "instantaneous" naked singularity with zero mass – *a naked singularity that exists for an instant and then it continues its electrical radiation to create negative mass*. If we set the mass function $f(u) = 0$ in (4.12), the metric will represent an 'instantaneous' naked singularity with zero mass [1], and at that stage the gravity will depend only on electrical charge, *i.e.* $\psi_2 = e^2/r^4$, and not on the mass of black hole. However, the Maxwell scalar is unaffected which leads to the proof of non-rotating part of Theorem 5.

The time taken between two consecutive radiation is so short that one cannot physically realize how quickly radiation takes place. Immediately, after the exhaustion of the Reissner-Nordstrom mass, if one still continues to radiate electrically with $e(r)$, there may be formation of new mass m_1^* (say). If this electrical radiation process continues further, the new mass will increase gradually as

$$\mathcal{M}^* = m_1^* + m_2^* + m_3^* + \dots \quad (4.15)$$

This new mass will never decrease. Since the mass increases, the temperature will also start increasing from zero. Then, the space-time metric will take the form

$$ds^2 = [1 + r^{-2}\{2r(\mathcal{M}^* - f(u)) + e^2\}] du^2$$

$$+2du dr - r^2 d\theta^2 - r^2 \sin^2 \theta d\phi^2. \quad (4.16)$$

This metric will describe a black hole if $f(u) - \mathcal{M}^* > e^2$, that is, when $f(u) > \mathcal{M}^* + e^2$. Thus, we have shown the changes in the total mass of Reissner-Nordstrom-Vaidya black hole in classical space-time metrics without affecting the Maxwell scalar ϕ_1 and the Vaidya mass function $f(u)$, for every electrical radiation during the Hawking evaporation process. We have also observed that, when $f(u) > \mathcal{M}^*$, the presence of Vaidya mass $f(u)$ in (4.16) can prevent the direct formation of negative mass naked singularity. Otherwise, when $f(u) < \mathcal{M}^*$, this metric may describe a ‘non-stationary’ negative mass naked singularity, which is different from the ‘stationary’ one discussed in [1]. The metric (4.16) can be written in Kerr-Schild ansatz in different backgrounds as

$$g_{ab}^{\text{NMV}} = g_{ab}^{\text{V}} + 2Q(r)\ell_a\ell_b,$$

where $Q(r) = (r\mathcal{M}^* + e^2/2)r^{-2}$ in Vaidya background, and

$$g_{ab}^{\text{NMV}} = g_{ab}^{\text{NM}} + 2Q(u, r)\ell_a\ell_b,$$

with $Q(u, r) = -f(u)r^{-1}$. These Kerr-Schild forms show that the metric (4.1) is a solution of Einstein’s field equations and completes the proof of non-rotating part of Theorem 2 as mention in the introduction. Here the metric tensor g_{ab}^{V} is non-rotating Vaidya null radiating metric, and g_{ab}^{NM} is that of the negative mass naked singularity with mass \mathcal{M}^* . The metric (4.16) leads to the proof of the first part of Theorem 7.

5 Changing mass of variable charged Kerr-Newman-Vaidya black hole

In order to show the Hawking’s radiation in the embedded Kerr-Newman-Vaidya black hole, we must have a different stress-energy tensor as the Kerr-Newman-Vaidya black hole with T_{ab} (3.2) does not have any direct Hawking’s radiation effects. The consideration of different forms of stress-energy-momentum tensor in the study of Hawking’s radiation effect in classical space-time metrics is followed from Boulware’s suggestion [3] mentioned in introduction above.

It is noted that the Kerr-Newman-Vaidya black hole, describing the Kerr-Newman black hole embedded into the rotating null radiating Vaidya, is quite different from the standard Kerr-Newman black hole. That is to say that, (i) the Kerr-Newman-Vaidya black hole is algebraically special in Petrov classification with the Weyl scalars $\psi_0 = \psi_1 = 0$, $\psi_2 = \psi_3 = \psi_4 \neq 0$, where as the standard Kerr-Newman black hole is Petrov type D with $\psi_0 = \psi_1 = \psi_3 = \psi_4 = 0$, $\psi_2 \neq 0$. (ii) the Kerr-Newman-Vaidya black hole possesses the total energy-momentum tensor (3.2), representing interaction of the null fluid $T_{ab}^{(n)}$ with the electromagnetic field $T_{ab}^{(E)}$, *i.e.*,

the charged null radiating fluid; however the energy-momentum tensor of the Kerr-Newman black hole is that of electromagnetic field $T_{ab}^{(E)}$, simply a charged black hole. (iii) the Kerr-Newman-Vaidya black hole is a non-stationary extension of the stationary Kerr-Newman black hole. Due to the above differences between the two black holes, it is worth to study the physical properties of the embedded black hole. Hence, it is hoped that the study of Hawking electrical radiation in the embedded Kerr-Newman-Vaidya black hole will certainly lead to a different physical feature than that of *non-embedded* Kerr-Newman black hole.

Thus, we consider the line element of a variable-charged Kerr-Newman-Vaidya black hole with the charge e varying with respect to the coordinate r in (3.1) as follows:

$$\begin{aligned}
ds^2 = & [1 - R^{-2}\{2r(M + f(u)) - e^2(r)\}] du^2 + 2du dr \\
& + 2aR^{-2}\{2r(M + f(u)) - e^2(r)\} \sin^2\theta du d\phi - 2a \sin^2\theta dr d\phi \\
& - R^2 d\theta^2 - \{(r^2 + a^2)^2 - \Delta^* a^2 \sin^2\theta\} R^{-2} \sin^2\theta d\phi^2,
\end{aligned} \tag{5.1}$$

where $\Delta^* = r^2 - 2r\{M + f(u)\} + a^2 + e^2(r)$. Then, with the variable charge $e(r)$, there will be a different energy-momentum tensor from that of the original Kerr-Newman-Vaidya metric (3.1). Accordingly, we have calculated the Ricci and Weyl scalars from the field equations with the functions $e(r)$ and found as follows:

$$\begin{aligned}
\phi_{11} = & \frac{1}{2R^2 R^2} \left\{ e^2(r) - 2r e(r) e(r)_{,r} \right\} \\
& + \frac{1}{4R^2} \left\{ e^2(r)_{,r} + e(r) e(r)_{,rr} \right\},
\end{aligned} \tag{5.2}$$

$$\begin{aligned}
\phi_{22} = & -\frac{r}{2R^2 R^2} \left\{ 2r f(u)_{,u} + a^2 \sin^2\theta f(u)_{,uu} \right\}, \\
\phi_{12} = & \frac{i a \sin\theta}{\sqrt{2} K \bar{R} R^2} f(u)_{,u},
\end{aligned}$$

$$\Lambda = -\frac{1}{12R^2} \left\{ e^2(r)_{,r} + e(r) e(r)_{,rr} \right\}, \tag{5.3}$$

$$\begin{aligned}
\psi_2 = & \frac{1}{\bar{R} \bar{R} R^2} \left[e^2(r) - R\{M + f(u)\} \right. \\
& \left. + \frac{1}{6} \bar{R} \bar{R} \left\{ e^2(r)_{,r} + e(r) e(r)_{,rr} \right\} - \bar{R} e(r) e(r)_{,r} \right],
\end{aligned} \tag{5.4}$$

$$\psi_3 = \frac{-i a \sin\theta}{2\sqrt{2} \bar{R} \bar{R} R^2} \left\{ (4r + \bar{R}) f(u)_{,u} \right\},$$

$$\psi_4 = \frac{a^2 r \sin^2\theta}{2\bar{R} \bar{R} R^2 R^2} \left\{ R^2 f(u)_{,uu} - 2r f(u)_{,u} \right\}.$$

From these we observe that the variable charge $e(r)$ has the effect only on the scalars ϕ_{11} , Λ and ψ_2 . However, the non-vanishing of ϕ_{22} , ϕ_{12} , ψ_3 and ψ_4 indicate the different physical feature of non-stationary Kerr-Newman-Vaidya black hole from the stationary standard Kerr-Newman. One important feature of the field equations corresponding to the metric (5.1) with $e(r)$ is that the expressions for ϕ_{11} and Λ do not involve the Vaidya mass function $f(u)$. So it suggests the possibility

to study the Hawking's electrical radiation in this non-stationary embedded black hole.

Now for an electromagnetic field, the Ricci scalar $\Lambda \equiv (1/24)g^{ab}R_{ab}$ in (5.3) has to vanish leading to the solution

$$e^2(r) = 2rm_1 + C \quad (5.5)$$

where m_1 and C are real constants. Substituting this $e^2(r)$ in (5.2) and (5.4), and after identifying the constant $C \equiv e^2$, we obtain

$$\phi_{11} = \frac{e^2}{2R^2R^2}, \quad (5.6)$$

$$\psi_2 = \frac{1}{R\bar{R}R^2} \left[e^2 - R\{M - m_1 + f(u)\} \right], \quad (5.7)$$

and ϕ_{12} , ϕ_{22} , ψ_3 and ψ_4 are remained unaffected as in (5.2) and (5.4) above. Accordingly, the Maxwell scalar $\phi_1 = \frac{1}{2}F_{ab}(\ell^a n^b + \bar{m}^a m^b)$ with the constant charge e becomes

$$\phi_1 = \frac{e}{\sqrt{2R\bar{R}}}, \quad (5.8)$$

which is the same Maxwell scalar of the Kerr-Newman-Vaidya metric (3.1) if once written out from (B.3) below. From the Weyl scalar (5.7) we have clearly seen a *change in the mass M* by some constant quantity m_1 for the first electrical radiation in the embedded black hole. Then the total mass of Kerr-Newman-Vaidya black hole becomes $(M - m_1) + f(u)$ and the metric takes the form

$$\begin{aligned} ds^2 = & [1 - R^{-2}\{2r(M - m_1 + f(u)) - e^2\}] du^2 + 2du dr \\ & + 2aR^{-2}\{2r(M - m_1 + f(u)) - e^2\} \sin^2\theta du d\phi - 2a \sin^2\theta dr d\phi \\ & - R^2 d\theta^2 - \{(r^2 + a^2)^2 - \Delta^* a^2 \sin^2\theta\} R^{-2} \sin^2\theta d\phi^2, \end{aligned} \quad (5.9)$$

where $\Delta^* = r^2 - 2r\{M - m_1 + f(u)\} + a^2 + e^2$. Since the Maxwell scalar (5.8) remains the same for the first Hawking radiation, we again consider the charge e to be a function of r for the second radiation in the metric (5.9) with the mass $M - m_1 + f(u)$. This process will certainly lead, by the Einstein-Maxwell field equations with the vanishing of Λ , to reduce another quantity m_2 (say) from the total mass, *i.e.* the mass becomes $M - (m_1 + m_2) + f(u)$ after the second electrical radiation. Here again, we observe that the Maxwell scalar ϕ_1 remains the same form and also there is no effect on the Vaidya mass $f(u)$ after the second radiation. Thus, if we consider n radiations, every time taking the charge e to be a function of r , the Einstein's field equations will imply that the total mass of the black hole takes the form

$$\mathcal{M} = M - (m_1 + m_2 + m_3 + m_4 + \dots + m_n) + f(u) \quad (5.10)$$

without affecting the mass function $f(u)$. Taking Hawking's radiation of charged black hole embedded in the rotating Vaidya null radiating space, one might expect that the mass M may be radiated away, just leaving the expression $M - (m_1 + m_2 + m_3 + m_4 + \dots + m_n)$ equivalent to the Planck mass of about 10^{-5} g and $f(u)$ untouched; that is, M may not be exactly equal to $(m_1 + m_2 + m_3 + m_4 + \dots + m_n)$, but has a difference of about a Planck-size mass. Otherwise, the mass M may be evaporated completely after continuous radiation, when $M = (m_1 + m_2 + m_3 + m_4 + \dots + m_n)$, just leaving the Vaidya mass $f(u)$ and the electric charge e only. Thus, this situation of the black hole without Kerr mass M may be seen in the form of a classical space-time metric as

$$\begin{aligned}
ds^2 = & [1 - R^{-2}\{2rf(u) - e^2\}] du^2 + 2du dr \\
& + 2aR^{-2}\{2rf(u) - e^2\} \sin^2\theta du d\phi - 2a \sin^2\theta dr d\phi \\
& - R^2 d\theta^2 - \{(r^2 + a^2)^2 - \Delta^* a^2 \sin^2\theta\} R^{-2} \sin^2\theta d\phi^2,
\end{aligned} \tag{5.11}$$

where $\Delta^* = r^2 - 2rf(u) + a^2 + e^2$. The Weyl scalar of this metric becomes

$$\psi_2 = \frac{1}{\bar{R} \bar{R} R^2} \{e^2 - R f(u)\}, \tag{5.12}$$

and ψ_3, ψ_4 are unaffected as (5.4). From (5.11) we know that the remaining metric is the rotating charged Vaidya null radiating black hole with $f(u) > a^2 + e^2$. The surface gravity for the solution at the horizon

$$r = r_{\pm} = f(u) \pm \sqrt{\{f(u)^2 - a^2 - e^2\}}, \tag{5.13}$$

is $\mathcal{K}_{\pm} = (1/2)R_{\pm}^{-2}\{r_{\mp} - r_{\pm}\}$. The Hawking's temperature on the horizons are $T_{\pm} = (1/4\pi)R_{\pm}^{-2}\{r_{\mp} - r_{\pm}\}$. The entropy and angular velocity of the horizon are found as

$$\begin{aligned}
\mathcal{S} &= 2\pi f(u) \left\{ f(u) + \sqrt{f(u)^2 - (a^2 + e^2)} \right\} - \pi e^2, \\
\Omega_{\pm} &= \frac{a\{2rf(u) - e^2\}}{(r^2 + a^2)^2} \Big|_{r=r_{\pm}}.
\end{aligned}$$

Thus, we may regard this left out remnant of the Hawking evaporation as the rotating charged Vaidya black hole. On the other hand, the metric (5.11) may be interpreted that the presence of Vaidya mass function $f(u)$ can avoid the formation of an 'instantaneous' naked singularity with zero mass. The formation of 'instantaneous' naked singularity with zero mass in *non-embedded* Reissner-Nordstrom and Kerr-Newman, black holes is unavoidable during Hawking's evaporation process, as shown in [1]. That is, if we set the mass function $f(u) = 0$, the metric (5.11) would certainly represent an 'instantaneous' naked singularity with zero mass, and at that stage gravity of the surface would depend only on electric charge, *i.e.* $\psi_2 = (e^2/\bar{R} \bar{R} R^2)$, and not on the mass of black hole. However, the Maxwell scalar ϕ_1 is unaffected. Thus, from (5.11) with $f(u) \neq 0$ it seems natural to refer

to the *rotating* charged Vaidya null radiating black hole as an ‘instantaneous’ black hole, during the Hawking’s evaporation process of Kerr-Newman-Vaidya black hole which proves the rotating part of Theorem 4.

The time taken between two consecutive radiations is supposed to be so short that one may not physically realize how quickly radiations take place. Immediately after the exhaustion of the Kerr-Newman mass, if one continues the remaining solution (5.11) to radiate electrically with $e(r)$, there may be a formation of a new mass m_1^* (say). If this electrical radiation process continues forever, the new mass will increase gradually as

$$\mathcal{M}^* = m_1^* + m_2^* + m_3^* + m_4^* + \dots . \quad (5.14)$$

However, it appears that this new mass will never decrease. Then, the space-time metric will take the following form

$$\begin{aligned} ds^2 = & [1 + R^{-2}\{2r(\mathcal{M}^* - f(u)) + e^2\}] du^2 + 2du dr \\ & + 2aR^{-2}\{2r(f(u) - \mathcal{M}^*) - e^2\} \sin^2\theta du d\phi - 2a \sin^2\theta dr d\phi \\ & - R^2 d\theta^2 - \{(r^2 + a^2)^2 - \Delta^* a^2 \sin^2\theta\} R^{-2} \sin^2\theta d\phi^2, \end{aligned} \quad (5.15)$$

where $\Delta^* = r^2 - 2r\{f(u) - \mathcal{M}^*\} + a^2 + e^2$. This metric will describe a black hole if $f(u) - \mathcal{M}^* > a^2 + e^2$, that is, when $f(u) > \mathcal{M}^* > a^2 + e^2$ for a particular value of u . Thus, we have shown the changes in the total mass of Kerr-Newman-Vaidya black hole in classical space-time metrics without affecting the Maxwell scalar and the Vaidya mass, for every electrical radiation during the Hawking evaporation process. We have also observed that, when $f(u) > \mathcal{M}^*$, the presence of Vaidya mass $f(u)$ in (5.15) can prevent the direct formation of negative mass naked singularity. Otherwise, when $f(u) < \mathcal{M}^*$, this metric may describe a ‘non-stationary’ negative mass naked singularity, which is different from the ‘stationary’ one discussed in [1]. The metric (5.15) can be expressed in Kerr-Schild ansatz in different backgrounds as

$$g_{ab}^{\text{NMV}} = g_{ab}^{\text{V}} + 2Q(r, \theta)\ell_a\ell_b, \quad (5.16)$$

where $Q(r, \theta) = (r\mathcal{M}^* + e^2/2)R^{-2}$, and

$$g_{ab}^{\text{NMV}} = g_{ab}^{\text{NM}} + 2Q(u, r, \theta)\ell_a\ell_b, \quad (5.17)$$

with $Q(u, r, \theta) = -rf(u)R^{-2}$. These Kerr-Schild forms show that the metric (5.15) is a solution of Einstein’s field equations. Here the metric tensor g_{ab}^{V} is rotating Vaidya null radiating black hole metric, and g_{ab}^{NM} is the metric describing the rotating negative mass naked singularity. The metric (5.15) leads to the proof of Theorem 7 for creating negative mass naked singularity of the rotating part, as mentioned in the introduction.

6 Conclusion

In this paper, by adopting Wang-Wu mass function we derive a class of embedded exact solutions of Einstein's field equations, namely non-rotating Reissner-Nordstrom-Vaidya and rotating Kerr-Newman-Vaidya solutions. The gravitational structure of the solutions are analyzed by observing the nature of the energy momentum tensor. The energy-momentum tensors (2.4) and (3.2) associated with the embedded solutions obey the energy conservation equations $T^{ab}{}_{;b} = 0$, showing that the space-time metrics (2.3) and (3.1) are exact solutions of field equations. These tensors show the interaction of Vaidya null fluid $T_{ab}^{(n)}$ with the electromagnetic field $T_{ab}^{(E)}$ – non-rotating in (2.4) and rotating in (3.2).

These embedded non-rotating Reissner-Nordstrom-Vaidya and rotating Kerr-Newman-Vaidya solutions can also be able to expressed in Kerr-Schild ansatze. The solution of Reissner-Nordstrom-Vaidya is of Petrov type D; and Kerr-Newman-Vaidya solution is algebraically special in Petrov classification of space-time metric. In embedded Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya black hole, the presence of Vaidya mass completely prevent the disappearance of black hole masses during the radiation process and thereby, the formation of “instantaneous” charged Vaidya black holes, non-rotating in (4.12) and rotating in (5.11). However, the non-embedded Reissner-Nordstrom and Kerr-Newman cannot prevent the disappearance of the total mass of the evaporating black holes.

It appears that (i) the changes in the mass of black holes, (ii) the formation of ‘instantaneous’ naked singularities with zero mass and (iii) the creation of ‘negative mass naked singularities’ in *non-embedded* Reissner-Nordstrom as well as Kerr-Newman black holes [1] are presumably the correct formulation in classical space-time metrics of the three possibilities of black hole evaporation suggested by Hawking and Israel [16]. However, the creation of ‘negative mass naked singularities’ may be a violation of Penrose’s cosmic censorship hypothesis [19]. The embedded black holes discussed here can be expressed in Kerr-Schild ansatze, accordingly their consequent negative mass naked singularities are also expressible in Kerr-Schild forms showing them as solution of Einstein’s field equations. It is also observed that once a charged black hole is embedded into some spaces, it will continue to embed into the same space forever through out its Hawking evaporation process. For example, Kerr-Newman black hole is embedded into the rotating Vaidya null radiating universe, it continues to embed as ‘instantaneous’ charged Vaidya black hole in (5.11) and embedded negative mass naked singularity as in (5.15). There Hawking’s radiation does not affect the Vaidya mass through out the evaporation process of Kerr-Newman mass. This means that the embedded negative mass naked singularities (5.15) possess the total energy momentum tensors (3.2), as the Kerr-Newman mass does not involved in these tensors (3.3) and (3.4), and the change in the mass due to continuous radiation does not affect them. Thus, it may be concluded that once a black hole is embedded into some spaces, it will continue to

embed forever without disturbing the nature of matters present in the backgrounds. Also since there is no effect on the Vaidya mass $f(u)$ during Hawking's evaporation process, it will always remain unaffected. That is to say that, unless some external forces apply to remove the mass $f(u)$ from the embedded space-time geometries, it will certainly continue to exist along with the electrically radiating black holes, rotating or non-rotating forever. If one accepts the Hawking continuous evaporation of charged black holes, the loss of the mass and creation of a new mass are the process of the continuous radiation. So, from the space-time metrics (5.9), (5.11), (5.15), it may also be concluded that once electrical radiation starts, it will continue to radiate forever describing the various stages of the life of radiating black holes. This establishes the proof of the rotating part of the Theorem 8 and simultaneously follows the non-rotating case of the theorem.

Also, we find from the above that the change in the mass of embedded black holes, takes place due to the Maxwell scalar ϕ_1 , remaining unchanged in the field equations during continuous radiation. So, if the Maxwell scalar ϕ_1 is absent from the space-time geometry, there will be no such electrical radiation, and consequently, there will be no observable changes in the mass of the black hole solution. Therefore, we cannot, theoretically, expect to observe such *relativistic change* in the mass of uncharged Schwarzschild as well as Kerr black holes. This suggests that these uncharged Schwarzschild as well as Kerr black holes will forever remain the same without changing their life styles.

From the study of Hawking's radiation above, it is also found that, as far as the embedded black holes are concerned, the Kerr-Newman black hole has relations with other rotating black holes, like the charged Vaidya black hole here. There the later ones are 'instantaneous' black holes of the respective embedded ones. It is observed that the classical space-time metrics discussed above would describe the possible life style of radiating embedded black holes at different stages during their continuous radiation. These *embedded* classical space-time metrics describing the changing life style of black holes are different from the *non-embedded* ones studied in [1] in various respects shown above. Here the study of these embedded solutions suggests the possibility that in an early universe there might be some black holes, which might have embedded into some other spaces possessing different matter fields with well-defined physical properties.

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Appendix A: Energy-momentum tensor and energy conservation equations

Here we present a rotating metric having a mass function $\hat{M}(u, r)$ given in Ref. [2] equation (6.4), which has been utilized in the derivation of the embedded Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya solutions.

$$\begin{aligned}
 ds^2 = & \{1 - 2r\hat{M}(u, r)R^{-2}\} du^2 + 2du dr \\
 & + 4a r \hat{M}(u, r)R^{-2} \sin^2 \theta du d\phi - 2a \sin^2 \theta dr d\phi \\
 & - R^2 d\theta^2 - \{(r^2 + a^2)^2 - \Delta a^2 \sin^2 \theta\} R^{-2} \sin^2 \theta d\phi^2, \quad (A.1)
 \end{aligned}$$

where $\Delta = r^2 - 2r\hat{M}(u, r) + a^2$ and $R^2 = r^2 + a^2 \cos^2 \theta$. When $a = 0$, this metric will reduce to a non-rotating canonical metric in Eddington-Finkelstein coordinate system,

$$ds^2 = \left\{ 1 - \frac{2}{r} \hat{M}(u, r) \right\} du^2 + 2du dr - r^2 (d\theta^2 + \sin^2 \theta d\phi^2). \quad (A.2)$$

By virtue of Einstein's field equations associated with the above line element (A.1) we obtain the total energy momentum tensor (EMT) as follows:

$$\begin{aligned}
 T_{ab} &= T_{ab}^{(n)} + T_{ab}^{(m)} \\
 &= \mu^* \ell_a \ell_b + 2\rho^* \ell_{(a} n_{b)} + 2p m_{(a} \bar{m}_{b)} + 2\omega \ell_{(a} \bar{m}_{b)} + 2\bar{\omega} \ell_{(a} m_{b)} \quad (A.3)
 \end{aligned}$$

where the EMTs for the rotating null fluid as well as that of the rotating matter are respectively given below:

$$\begin{aligned}
 T_{ab}^{(n)} &= \mu^* \ell_a \ell_b + 2\omega \ell_{(a} \bar{m}_{b)} + 2\bar{\omega} \ell_{(a} m_{b)} \\
 T_{ab}^{(m)} &= 2(\rho^* + p) \ell_{(a} n_{b)} - p g_{ab}, \quad (A.4)
 \end{aligned}$$

where $\bar{\omega}$ is the complex conjugate of ω referred to as a rotational density. When $a = 0$, the rotational density ω will be zero, and then these EMTs are similar to those introduced by Husain [27] in the case of non-rotating fluid. Here the quantities μ^* , ρ^* , p and ω^* in (A.3) are found as

$$\begin{aligned}
 K\mu^* &= -\frac{1}{R^2 R^2} \left\{ 2r^2 \hat{M}(u, r)_{,u} + a^2 r \sin^2 \theta \hat{M}(u, r)_{,uu} \right\} \\
 K\rho^* &= \frac{2r^2}{R^2 R^2} \hat{M}(u, r)_{,r} \\
 Kp &= \frac{1}{R^2 R^2} \left\{ 2r^2 \hat{M}(u, r)_{,r} - R^2 \left(2\hat{M}(u, r)_{,r} + r \hat{M}(u, r)_{,rr} \right) \right\} \\
 K\omega &= -\frac{i a \sin \theta}{\sqrt{2} R^2 R^2} \left\{ R \hat{M}(u, r)_{,u} - r \bar{R} \hat{M}(u, r)_{,ru} \right\} \quad (A.5)
 \end{aligned}$$

with the universal constant $K = 8\pi G/c^4$. The trace of $T_{ab}^{(n)}$ is zero and that of $T_{ab}^{(m)}$ is $T = 2(\rho^* - p)$. The line element (A.1) contains many known spherically rotating axisymmetric solutions.

The energy conservation equations for the energy-momentum tensor (A.2) $T^{ab}{}_{;b} = 0$ can, in general, be transcribed in Newman-Penrose spin coefficient formalism as follows

$$D\rho^* = \rho^*(\rho + \bar{\rho}) + p(\rho + \bar{\rho}) + \omega\bar{\kappa} + \bar{\omega}\kappa \quad (\text{A.6})$$

$$\begin{aligned} D\mu^* + \nabla\rho^* + \bar{\delta}\omega + \delta\bar{\omega} &= \mu^*\{(\rho + \bar{\rho}) - 2(\epsilon + \bar{\epsilon})\} - \rho^*(\mu + \bar{\mu}) - p(\mu + \bar{\mu}) \\ &\quad - \omega(2\pi + 2\bar{\beta} - \bar{\tau}) - \bar{\omega}(2\bar{\pi} + 2\beta - \tau) \end{aligned} \quad (\text{A.7})$$

$$\begin{aligned} D\bar{\omega} + \bar{\delta}p &= \mu^*\bar{\kappa} + \rho^*(\bar{\tau} - \pi) + p(\bar{\tau} - \pi) + \omega\bar{\sigma} \\ &\quad - \bar{\omega}(2\epsilon - 2\bar{\rho} - \rho) \end{aligned} \quad (\text{A.8})$$

where κ, τ, π etc are spin coefficients and the derivative operators are defined as

$$D \equiv \ell^a \partial_a, \quad \nabla \equiv n^a \partial_a, \quad \delta \equiv m^a \partial_a, \quad \bar{\delta} \equiv \bar{m}^a \partial_a. \quad (\text{A.9})$$

These (A.6-A.8) are general equations for the energy conservation laws associated with any energy-momentum tensor of the type (A.3). An energy-momentum tensor (A.3) for any rotating solution has to satisfy these energy conservation equations; then the metric can to be called an exact solution of Einstein's equations.

Appendix B: NP spin coefficients for the Kerr-Newman-Vaidya black hole

For the analysis of the embedded Kerr-Newman-Vaidya black hole (3.1), we cite the complex null tetrad vectors for the line element as follows

$$\begin{aligned} \ell_a &= \delta_a^1 - a \sin^2 \theta \delta_a^4, \\ n_a &= \frac{\Delta}{2R^2} \delta_a^1 + \delta_a^2 - \frac{\Delta}{2R^2} a \sin^2 \theta \delta_a^4, \\ m_a &= -\frac{1}{\sqrt{2R}} \left\{ -ia \sin \theta \delta_a^1 + R^2 \delta_a^3 + i(r^2 + a^2) \sin \theta \delta_a^4 \right\}, \end{aligned} \quad (\text{B.1})$$

where $\Delta = r^2 - 2r\{M + f(u)\} + a^2 + e^2$. Using these null tetrad vectors, we solve the Cartan's first structure equations, transcribed in NP formalism by McIntosh and Hickman [19], for the spin coefficients which are found as follows

$$\begin{aligned} \kappa &= \epsilon = \lambda = \sigma = 0, \\ \rho &= -\frac{1}{R}, \quad \mu = -\frac{\Delta}{2RR^2}, \\ \alpha &= \frac{(2ai - R \cos \theta)}{2\sqrt{2R} R \sin \theta}, \quad \beta = \frac{\cot \theta}{2\sqrt{2R}}, \\ \pi &= \frac{ia \sin \theta}{\sqrt{2R} R}, \quad \tau = -\frac{ia \sin \theta}{\sqrt{2R^2}}, \end{aligned} \quad (\text{B.2})$$

$$\gamma = \frac{1}{2\bar{R}R^2} [\{r - m - f(u)\}\bar{R} - \Delta].$$

The tetrad components of Ricci and Weyl tensors, known as Ricci and Weyl scalars respectively are obtained from the Cartan's second structure equations. The Ricci scalars are as follows

$$\begin{aligned}\phi_{00} &= \phi_{01} = \phi_{10} = \phi_{20} = \phi_{02} = \Lambda = 0 \\ \phi_{11} &= \frac{e^2}{2R^2\bar{R}^2} \\ \phi_{12} &= \frac{ia\sin\theta}{2\sqrt{2}R^2\bar{R}}f(u)_{,u} \\ \phi_{22} &= -\frac{r}{2K R^2\bar{R}}\{2rf(u)_{,u} + a^2\sin^2\theta f(u)_{,uu}\}.\end{aligned}\tag{B.3}$$

Here it is found that the Ricci scalar $\Lambda = (1/24)g^{ab}R_{ab}$ vanishes for the rotating solution (3.1). This shows the characteristic feature of Kerr-Newman-Vaidya that it is an electrically charged black hole. The Maxwell scalar ϕ_1 can be obtained from the Ricci scalar ϕ_{11} using the relation $\phi_{11} = \phi_1\bar{\phi}_1$ for electromagnetic field F_{ab} [17]. The Weyl scalars become

$$\begin{aligned}\psi_0 &= \psi_1 = 0 \\ \psi_2 &= -\frac{1}{\bar{R}\bar{R}R^2}\left[R\{m + f(u)\} - e^2\right] \\ \psi_3 &= -\frac{ia\sin\theta}{2\sqrt{2}\bar{R}\bar{R}R^2}\left\{(4r + \bar{R})f(u)_{,u}\right\} \\ \psi_4 &= \frac{a^2r\sin^2\theta}{2\bar{R}\bar{R}R^2R^2}\left\{R^2f(u)_{,uu} - 2rf(u)_{,u}\right\}.\end{aligned}\tag{B.4}$$

The vanishing of Weyl scalars ψ_0 and ψ_1 indicates that the Weyl curvature tensor of the solution is algebraically special type II in Petrov classification with a repeated principal null congruence ℓ_a (2.4), which is geodesic ($\kappa = \epsilon = 0$), shear free ($\sigma = 0$), expanding [$2\hat{\theta} \equiv \ell^a_{;a} = -(\rho + \bar{\rho})$] as well as non-zero twist [$4\hat{\omega}^2 \equiv 2\ell_{[a;b]}\ell^{a;b} = -(\rho - \bar{\rho})$]. Here $\kappa, \epsilon, \sigma, \rho$ are NP spin coefficients given in (B.2). This completes the proof of the Theorem 2 stated in the introduction above. For future use we also introduce the Ricci and Riemann Curvature invariants for the metric (3.1) as

$$\begin{aligned}R_{ab}R^{ab} &= \frac{4e^4}{R^2\bar{R}^2R^2R^2} \\ R_{abcd}R^{abcd} &= 24\left[\frac{(R\{m + f(u)\} - e^2)^2}{R^2\bar{R}^2\bar{R}\bar{R}R^2} + \frac{(\bar{R}\{m + f(u)\} - e^2)^2}{R^2\bar{R}^2RRR^2}\right] + \frac{20e^4}{R^2\bar{R}^2R^2R^2}.\end{aligned}\tag{B.5}$$

Here we shall verify that the components of T^{ab} with the quantities μ^*, ρ^*, p and ω given in (3.5a-c) for the rotating Kerr-Newman-Vaidya metric satisfy the

conservation equations (A.6-A.8). The directional derivative operators (A.9) along the null vectors (B.1) are given as follows:

$$\begin{aligned}
D &= \partial_r, \\
\nabla &= \frac{1}{R^2} \{ (r^2 + a^2) \partial_u - \frac{\Delta}{2} \partial_r + a \partial_\phi \}, \\
\delta &= \frac{1}{\sqrt{2}R} \{ ia \sin \theta \partial_u + \partial_\theta + \frac{i}{\sin \theta} \partial_\phi \}, \\
\bar{\delta} &= \frac{1}{\sqrt{2}R} \{ -ia \sin \theta \partial_u + \partial_\theta - \frac{i}{\sin \theta} \partial_\phi \}.
\end{aligned} \tag{B.6}$$

It is to say that the equations (A.6) and (A.8) are comparatively easier to verify than (A.7). Now, by virtue of (3.5) and (B.2), the left side of (A.7) is found as

$$\begin{aligned}
&D\mu^* + \nabla\rho^* + \bar{\delta}\omega + \delta\bar{\omega} \\
&= \frac{2re^2\Delta}{KR^2R^2R^2R^2} + \frac{2r^2a^2\sin^2\theta}{KR^2R^2R^2R^2} f(u)_{,uu} + \frac{2r}{KR^2R^2R^2R^2} \{ 2r^2 - a^2 \cos^2\theta \} f(u)_{,u} \\
&\quad + \frac{rf(u)_{,u}}{KR^2R^2R^2R^2} \{ a^2 \sin^2\theta (7a^2 \cos^2\theta - r^2) \},
\end{aligned} \tag{B.7}$$

which can be shown equal to the right side of (A.7), by using spin coefficients (B.2). It indicates the fact that the energy-momentum tensor (3.2) satisfies the conservation equation (3.6). Similarly, one can verify the conservation laws for the case of the non-rotating Reissner-Nordstrom-Vaidya metric (2.3) having T_{ab} (2.7), when $a = 0$. This shows the fact that the Reissner-Nordstrom-Vaidya and Kerr-Newman-Vaidya are solutions of Einstein's field equations.

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