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Search for star-planet interaction

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Abstract. We analyse the chromospherical activity of stars with extrasolar planets and search for a possible correlation between the equivalent width of the core of Ca II K line and orbital parameters of the planet. We found a statistically significant evidence that the equivalent width of the Ca II K line reversal, which originates in the stellar chromosphere depends on the orbital period $P_{\rm orb}$ of the exoplanet. Planets orbiting stars with $T_{\rm eff} < 5\,500$ K and with $P_{\rm orb} < 20$ days generally have much stronger emission than planets at similar temperatures but at longer orbital periods. $P_{\rm orb} = 20$ days marks a sudden change in behaviour, which might be associated with a qualitative change in the star-planet interaction.

Keywords. Ca II K line, exoplanet, star-planet interaction.

1. Introduction

The question of possible existence of star-planet interaction is currently studied in many ways. Based on the observations in the optical region Shkolnik et al. (2005, 2008) discovered the planetary induced variability in the cores of Ca II H & K, Hα and Ca II IR triplet in a few planet hosting stars. Knutson et al. (2010) found a correlation between the chromospheric activity of the star and presence of the stratosphere on the planet. Consequently, Hartman (2010) found a correlation between the surface gravity of Hot Jupiters and the stellar activity. Recently Canto Martins et al. (2011) searched for correlation between planetary parameters and the $\log R'_{\rm HK}$ parameter but didn't reveal any convincing proof for such a phenomenon.

2. Observation & Statistical Analysis

We used FEROS instrument on 2.2 ESO/MPG telescope to obtain spectra of several stars (HD 179949, HD 212301, HD 149143 and Wasp-18) with close-in exoplanet. We also used the publicly available spectra from HIRES spectrograph archive. Subsequently we measured the equivalent width of the central reversal in the core of Ca II K.

In the first case we divided our data sample into two groups according to the semi-major axis ($a \leqslant 0.15$ and $a > 0.15$ AU). Figure 1 (left-top) shows the dependence of equivalent width on the effective temperature of the star. Subsequently, we performed two statistical tests – Student's-t test and Kolmogorov-Smirnov test to determine whether the two groups originate from the same population. The resulting probability is a function of temperature and is plotted in the lower part of Figure 1. The tests show that the difference between the two samples is significant for $T_{\rm eff} \leqslant 5\,500$ K. It means that stars with lower temperature and with planets on closer orbits show more activity as measured in the core of Ca II K line.

In the second case we group the data according to the effective temperature of the parent star ($T_{\rm eff} \leqslant 5\,500$ K and $T_{\rm eff} > 5\,500$ K) and plot the equivalent width of the Ca II K line reversal as a function of the orbital period (Figure 1–right).

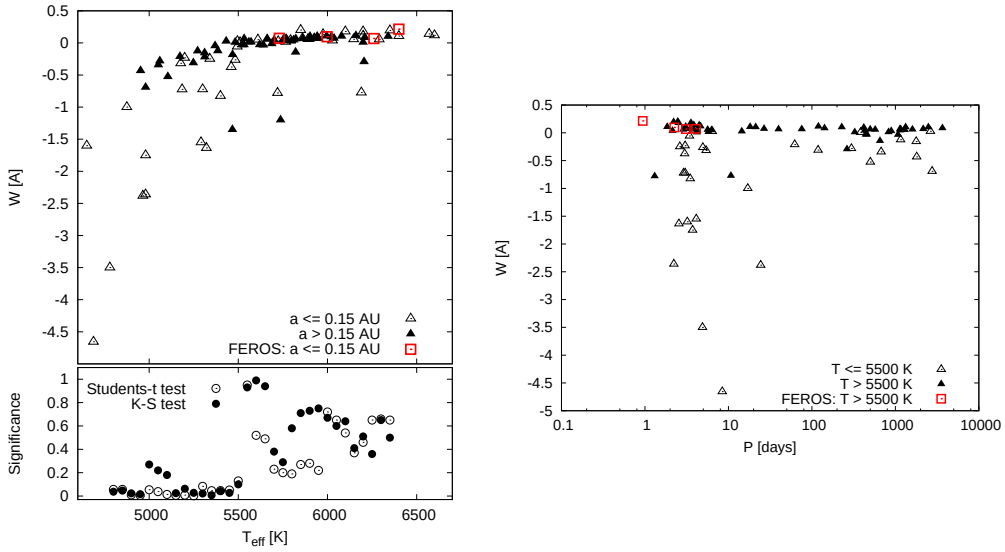

Figure 1. Left *Top*: Dependence of the equivalent width of Ca II K reversal on the temperature of the parent star. Empty triangles are exoplanetary systems with $a \leqslant 0.15$ AU, full triangles are systems with $a > 0.15$ AU. *Bottom*: Statistical Student's-t test (empty circles) and Kolmogorov-Smirnov test (full circles). Red squares are data from FEROS. **Right** Dependence of the equivalent width of Ca II K on the orbital period. Empty triangles are exoplanetary systems with $T \leqslant 5\,500$ K, full triangles are systems with $T > 5\,500$ K and red squares are data from FEROS.

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